



HORIZON2020 Marie Skłodowska-Curie actions

Numerical Relativity beyond General Relativity

Benasque, June 2018

Waveforms in Scalar-Tensor Theory of Gravity - Part I Anna Heffernan Collaborators: Laura Bernard, Ryan Lang, Cliff Will.





Scalar-Tensor Gravity

- * Alternate theory of gravity (ATG)
- Variable gravitational "constant" -> scalar field

$$S = \frac{1}{16\pi} \int \left[\phi R - \frac{1}{\phi} \omega(\phi) g^{\alpha\beta} \partial_{\alpha} \phi \partial_{\beta} \phi \right] \sqrt{-g} d^4 x + \int \mathcal{L}\left(m, g_{\alpha\beta} \sqrt{-g} d^4 x\right) \\ \Rightarrow T^{\alpha\beta}_{,\beta} = 0$$

- Obeys Einstein's Equivalence principle
- Violates Strong Equivalence principle

Internal Structure affects motion and GW emission

Wave equations via the Landau-Lifshitz formalism

$$\Box_{\eta}\tilde{h}^{\alpha\beta} = -16\pi\tau^{\alpha\beta}, \qquad \Box_{\eta}\Psi = -8\pi\tau_s.$$

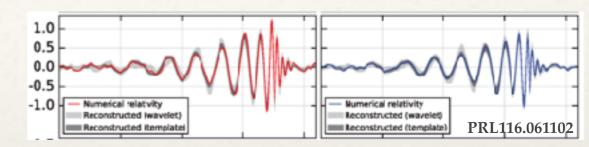
Motivation

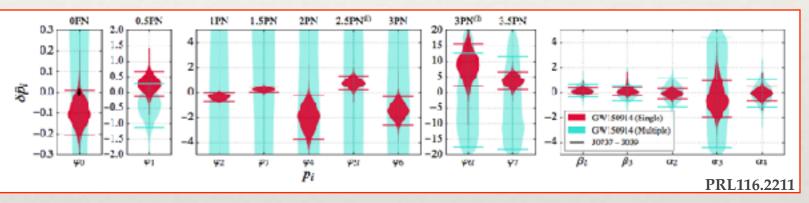
★ Seek to verify / constrain / discard ATG's → New tool: Gravitational Wave Astronomy



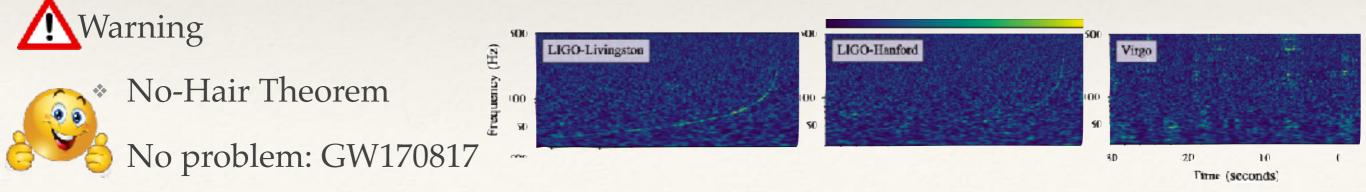
* Current Testing: Agnostic (See talks: Del Pozzo, Sennett)







- * Why Scalar Tensor?
 - * One of the simplest variations of GR
 - * Encapsulates some of F(R) and Superstring theories



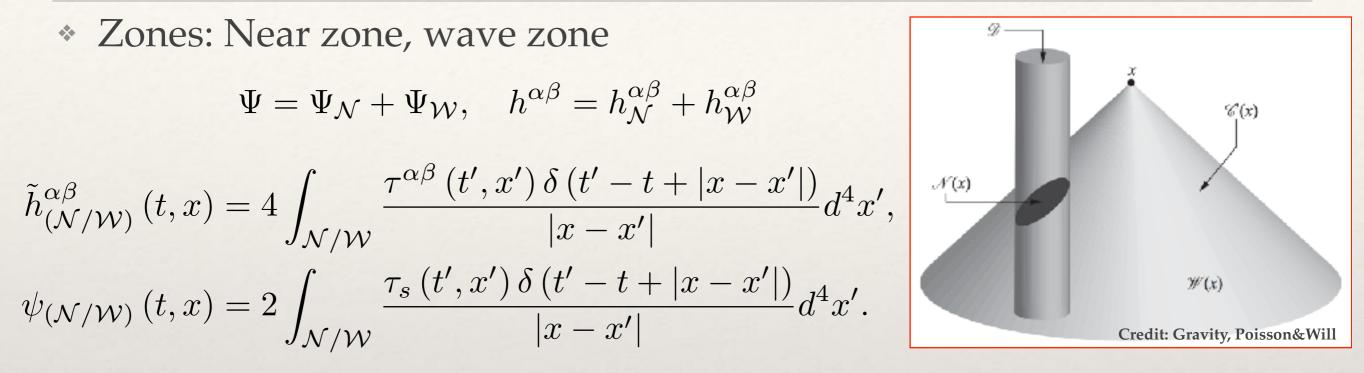
To date ...

- * Necessary ingredients for 2PN waveforms:
- ✤ Equations of Motion +0.5PN higher 3PN * 2.5PN EOM (Mirshekari & Will, 2013) EOM Tensor and scalar fields (and resulting energy fluxes) to 2PN 2PN Tensor gravitational waves and tensorial energy flux (Lang, 2014) 2.5PN * 1.5PN Scalar gravitational waves and 1PN scalar energy flux (Lang, 2015) Ψ Ready to use waveforms, incomplete 2PN (Sennett et al., 2016) Culprit: Non-vanishing scalar dipole moment $\Psi = \Psi_{-1/2} + \Psi_0 + \Psi_{1/2} + \Psi_1 + \Psi_{3/2}, \Rightarrow \text{We require } \Psi_{n+1/2} \text{ for } \dot{E}_n,$

 $\dot{E}_S \propto \dot{\Psi}^2 \Rightarrow \dot{E} = \dot{E}_{-1} + \dot{E}_0 + \dot{E}_{1/2} + \dot{E}_1.$ \Rightarrow We require EOM_(n+1) for $\Psi_{n+1/2}.$

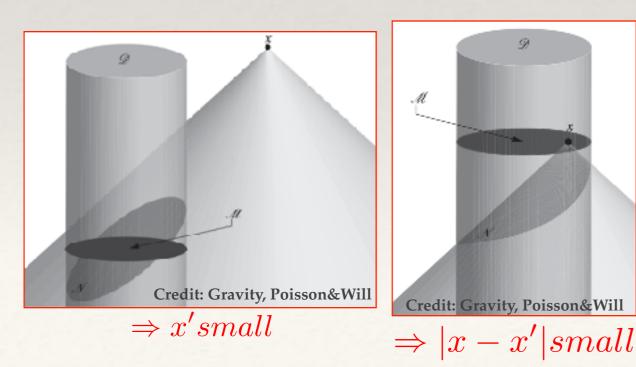
* 3PN EOM (Laura Bernard, 2018 - paper 1 of 2)

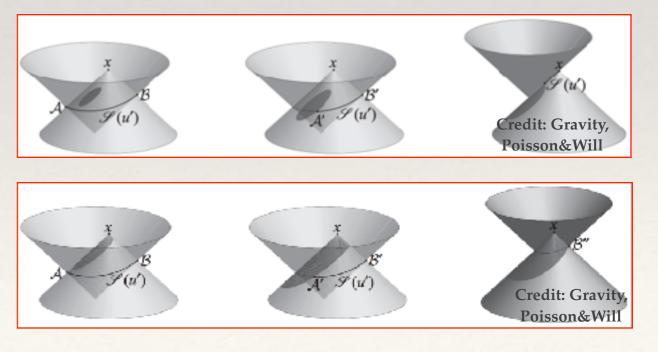
Direct Integration of Relaxed Einstein Equations



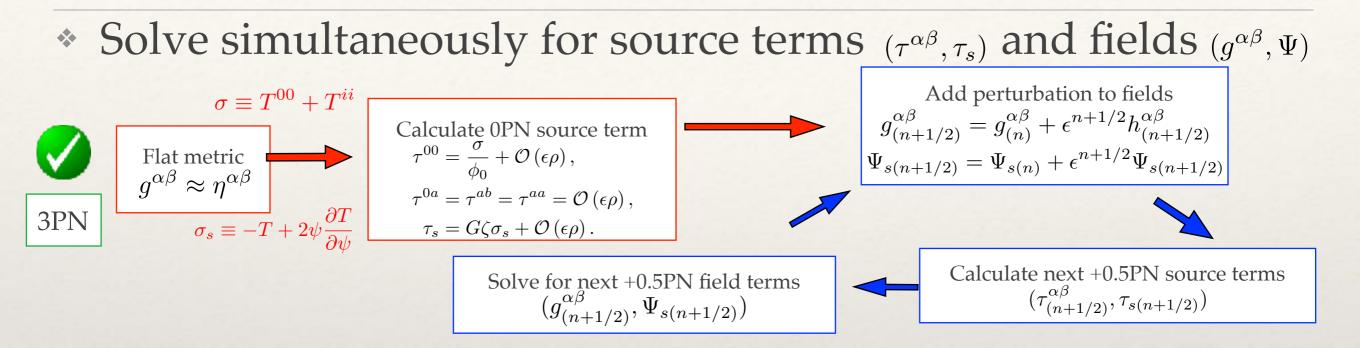
Near zone field

* Far zone field





Direct Integration of Relaxed Einstein Equations



* Use source terms to calculate multipole moments

$$\mathcal{I}^Q_{(s)} \equiv \int \tau_{(s)} x^Q d^3 x$$

 $G_{1}, G_{1s}, G_{2}, G_{2s}, G_{3}, G_{3s}, G_{4}, G_{5}, G_{6s}, v.G_{7}, H, H_{s}, H^{s}, H^{s}, U^{3}, U_{s}^{3}, U_{s}^{2}, v^{2} U_{s}^{2}, v^{2} U_{s}^{2}, v^{a} v^{b} P_{2ab}, v^{a} v^{b} P_{2ab}, v^{2} v.V, v.V_{2}, \overset{(a)}{Y}, \overset{(c)}{Y}, v^{2} \chi, \chi_{1}, \chi_{2}, v.\chi, v^{2} \chi_{s}, \chi^{s}_{2}, \chi^{s}_{2s}, \Sigma_{s}(\chi), \Sigma_{$

 $a_{s} G_{2s}, a_{s} G_{3s}, a_{s} G_{6s}, a_{s} H^{s}, a_{s} H^{s}, a_{s} H^{s}, a_{s} H^{s}, a_{s} H^{s}, a_{s} U_{s}^{3}, a_{s} U_{s}^{3}, a_{s} U_{s}^{3}, a_{s} U_{s}^{2}, b_{s} V^{2} U_{s}^{2}, a_{s} V^{2} U_{s}, a_{s} V^{2} X_{s}, a_{s} X^{s}_{2s}, a_{s} X^{s}_{2s}, a_{s} \Sigma_{s}(X), a_{s} \Sigma_{s}(X), a_{s} \Sigma_{s}(\Phi_{1}), a_{s} \Sigma_{s}(\Phi_{2}), a_{s} X^{s}_{s}(\Phi_{2}), a_{s} X^{s}_{s}(\Phi_{2}), a_{s} \Sigma_{s}(\Phi_{2}), a_{s} V^{2}\Phi_{2}, a_$

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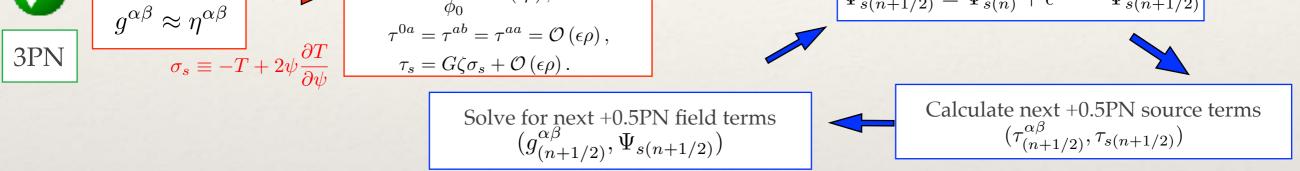
t_{s,F2}

 $p_{s} G_{1}, p_{s} G_{2}, p_{s} G_{5}, p_{s} U^{3}, p_{s} U_{s}^{3}, p_{s} V_{s}^{3}, p_{s} V_{s}^{2}, p_{s} \tilde{\chi}_{2}^{s}, p_{s} \tilde{\chi}_{s}^{s}, p_{s} \Sigma_{s}(\chi^{a}), p_{s} \Sigma_{s}(\Phi^{s}_{2}), p_{s} \Sigma_{s}(\Phi^{s}_{2}), p_{s} \Sigma_{s}(U_{s}^{2}), p_{s} \Sigma_{s}^{2}), p_{s} \Sigma_{s}^{2}(U_{s}^{2}), p_{s} \Sigma_{s}^{2}}, p_{s}^{2})$

 $\rho U^{2} U_{s} a_{s}, \rho_{s} U^{2} v^{2}, \rho_{s} U^{3}, \rho_{s} U^{2} U_{s}, \rho \Phi_{2} U_{s} a_{s}, \rho_{s} \Phi_{2} v^{2}, \rho_{s} \Phi_{2} U_{s}, \rho \Phi_{2}^{s} U_{s} a_{s}, \rho_{s} \Phi_{2}^{s} v^{2}, \rho_$

Direct Integration of Relaxed Einstein Equations





Use source terms to calculate multipole moments

$$\mathcal{I}^Q_{(s)} \equiv \int \tau_{(s)} x^Q d^3 x$$

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 $\begin{array}{c} u_{1}^{(1)}(i,j_{k}^{(0)},(i,j_{k}),u_{k},u_{k},w_{k},u_{k$

* Terms like \mathcal{I} require 3PN equations of motion

Tricky in Will, Wiseman & Pati A Solution: Laura Bernard!

And now to Laura