Axion searches

TAE 2018

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based mostly on ...

1801.08127

New experimental approaches in the search for axion-like particles

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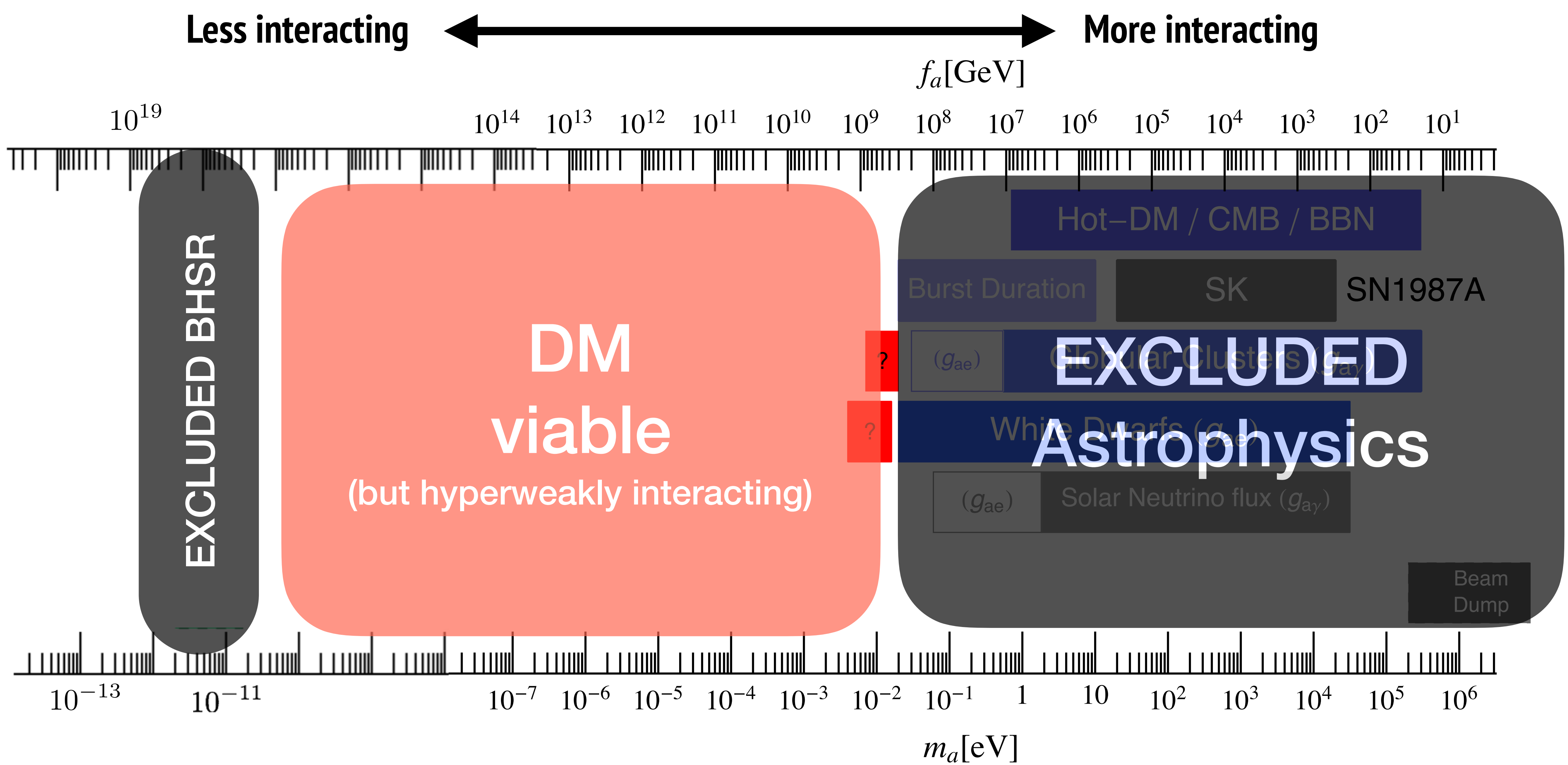
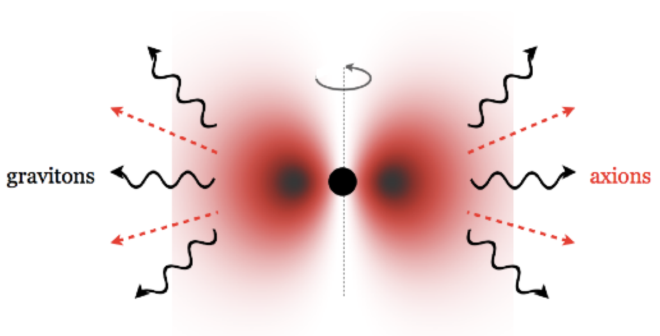
May 8, 2018

Abstract

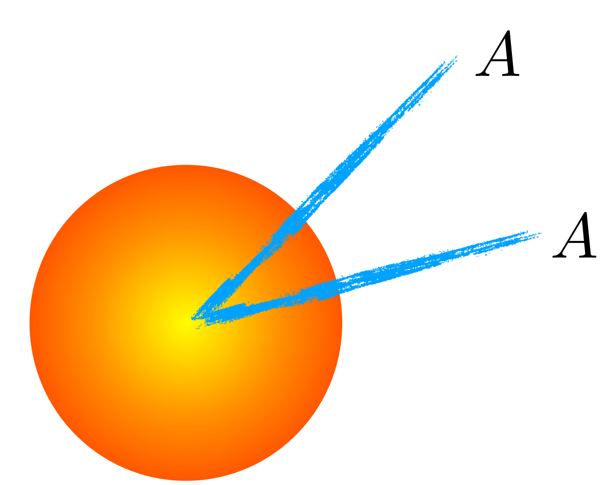
Axions and other very light axion-like particles appear in many extensions of the Standard Model, and are leading candidates to compose part or all of the missing matter of the Universe. They also appear in models of inflation, dark radiation, or even dark energy, and could solve some long-standing astrophysical anomalies. The physics case of these particles has been considerably developed in recent years, and there are now useful guidelines and powerful motivations to attempt experimental detection. Admittedly, the lack of a positive signal of new physics at the high energy frontier, and in underground detectors searching for weakly interacting massive particles, is also contributing to the increase of interest in axion searches. The experimental landscape is rapidly evolving, with many novel detection concepts and new experimental proposals. An updated account of those initiatives is lacking in the literature. In this review we attempt to provide such an update. We will focus on the new experimental approaches and their complementarity, but will also review the most relevant recent results from the consolidated strategies and the prospects of new generation experiments under consideration in the field. We will also briefly review the latest developments of the theory, cosmology and astrophysics of axions and we will discuss the prospects to probe a large fraction of relevant parameter space in the coming decade.

what do we know about fA





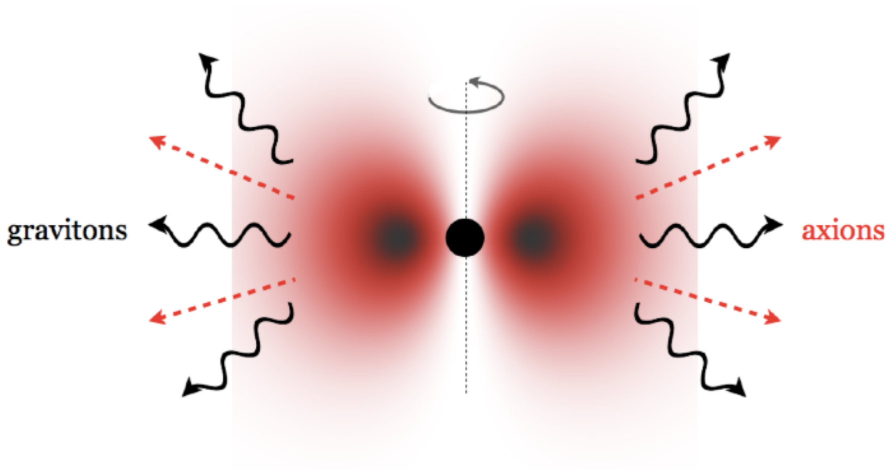
Black hole spin radiated



Stellar evolution accelerated*

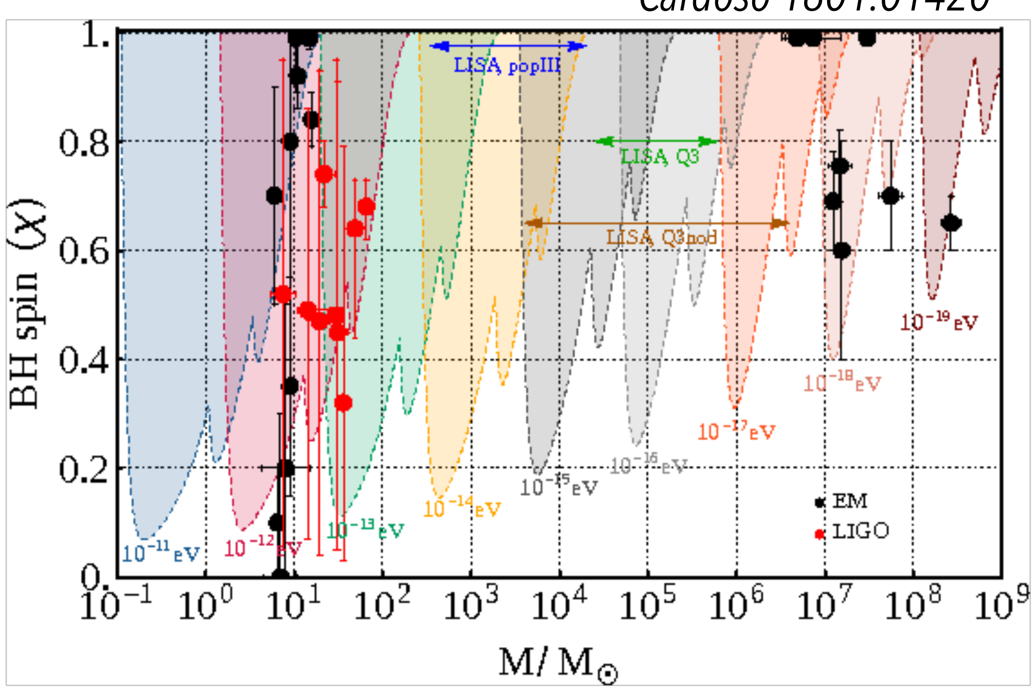
BH superradiance (indep. of DM hypothesis)

-Black-hole super radiance $R_{\rm BH} \sim 1/m_A$ **BHs with spin can radiate it into axions quite efficiently,**



- we observe spin BH's so we can put constraints…

Cardoso 1801.01420



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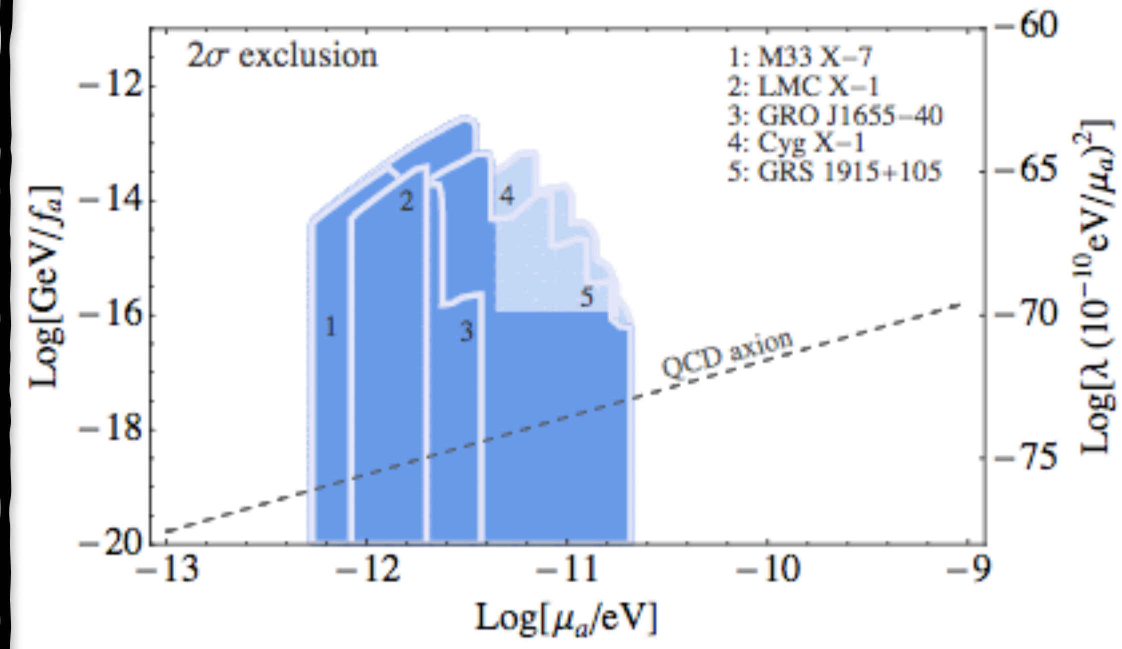
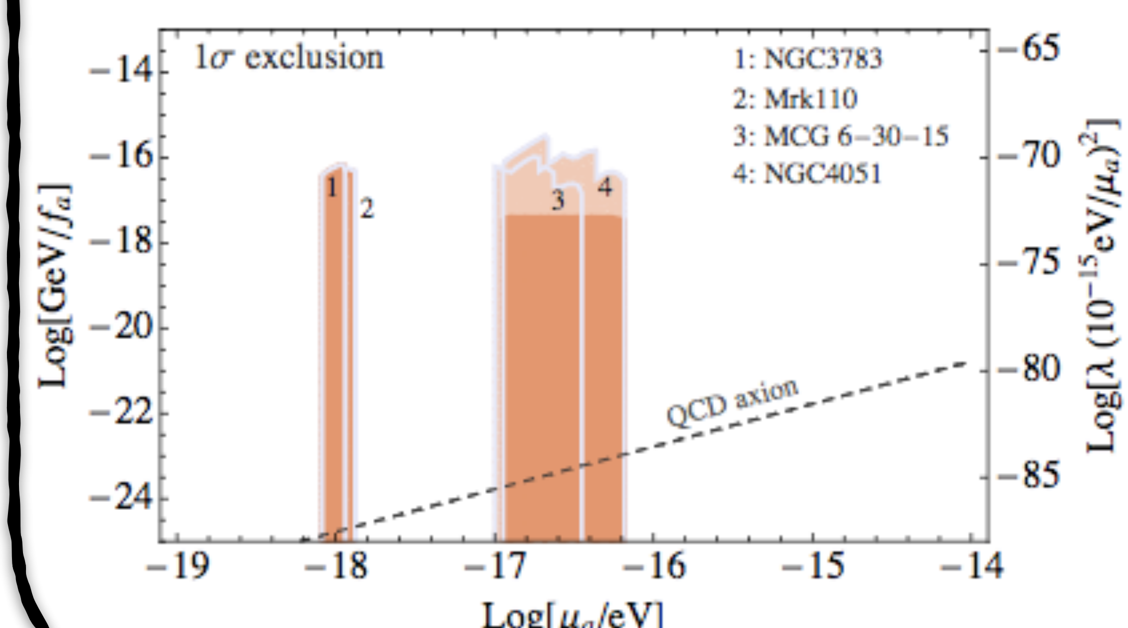
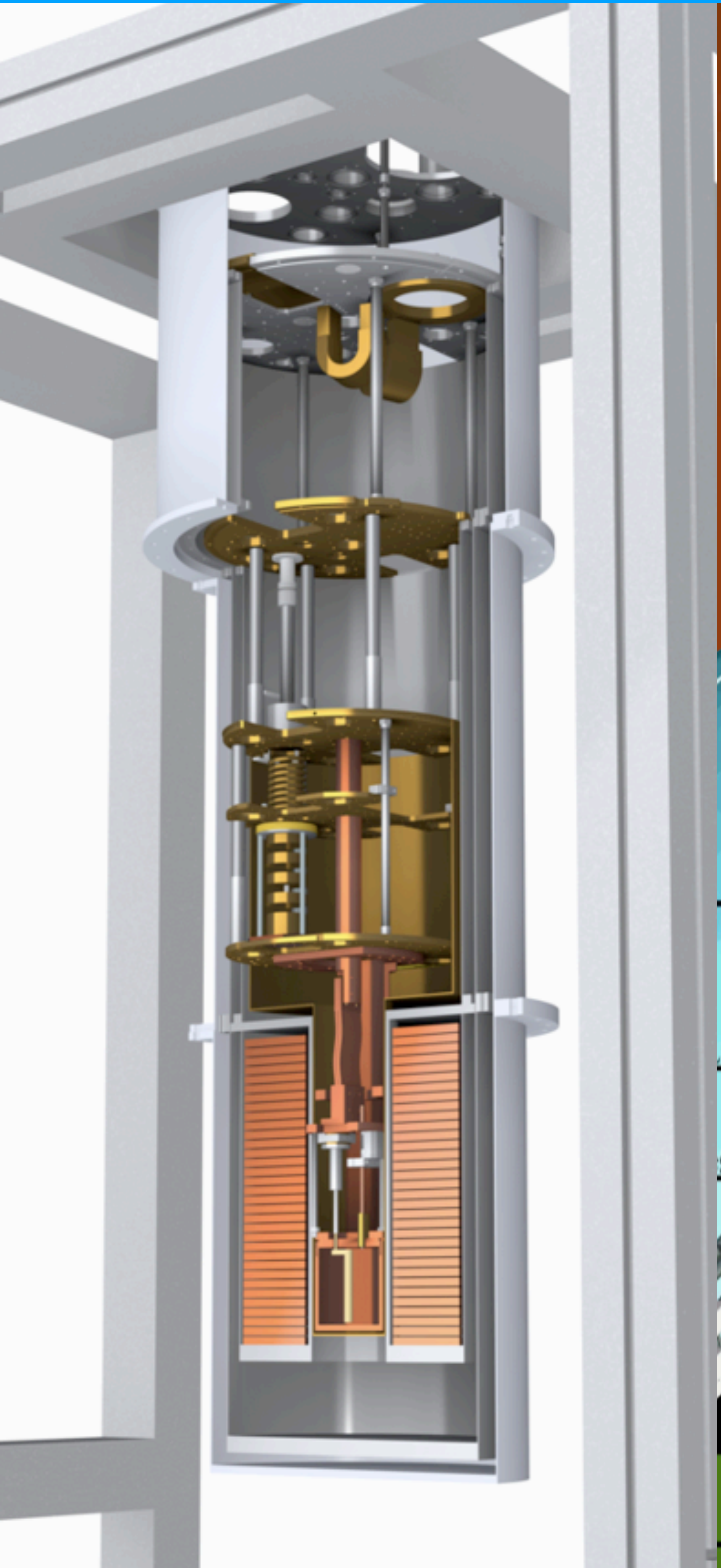


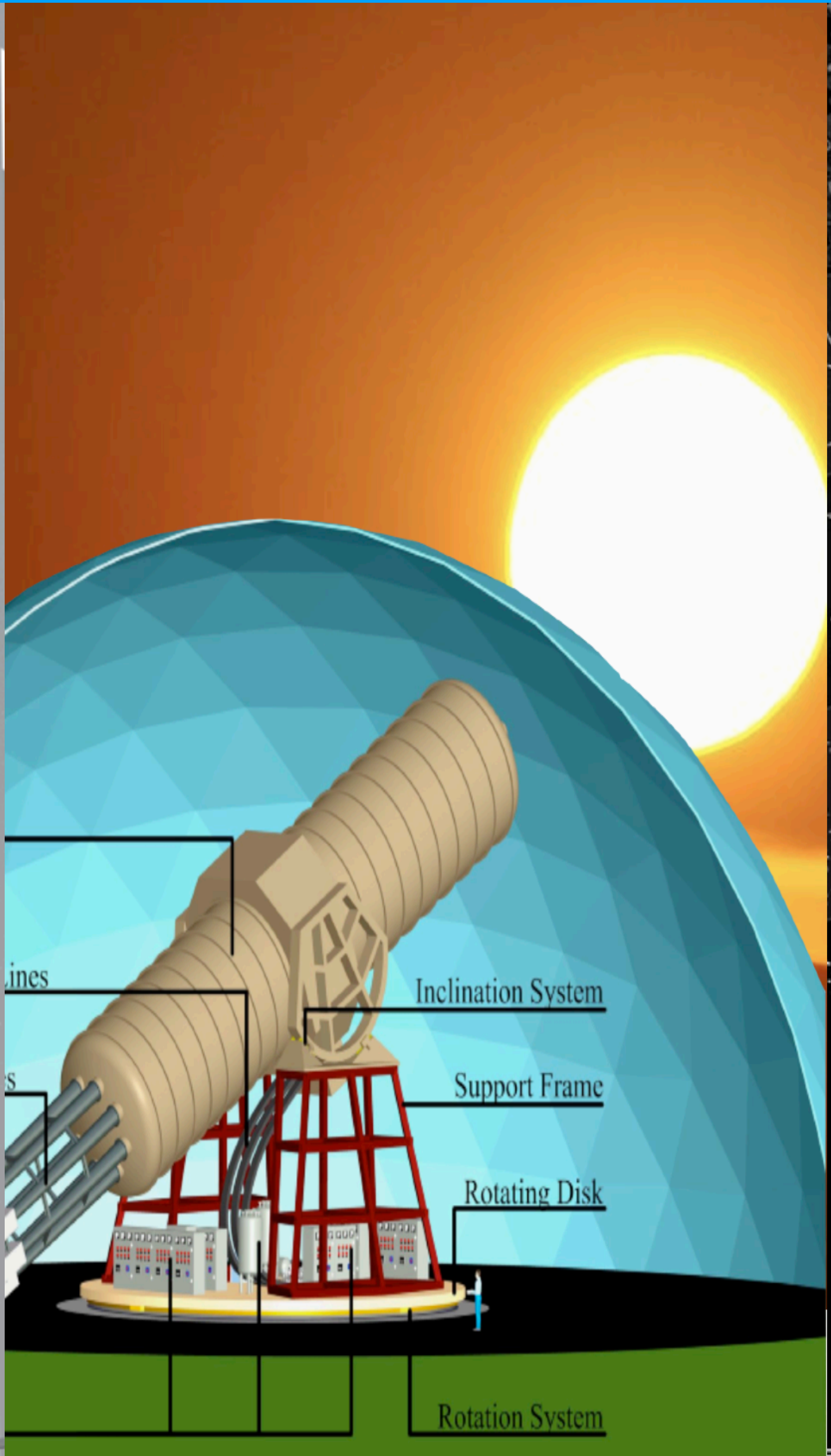
(a)



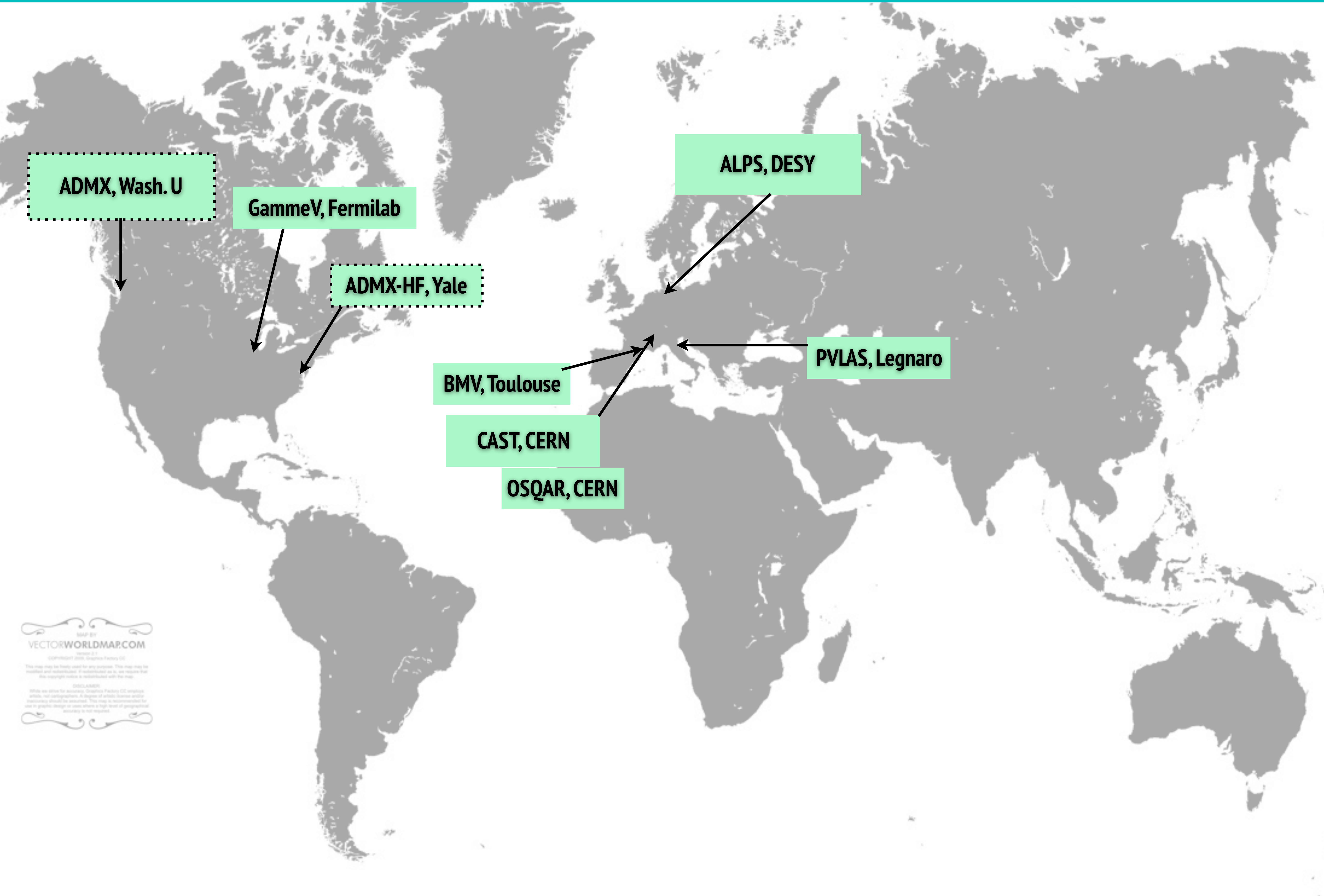
(b)

Lab	Stars	Cosmos

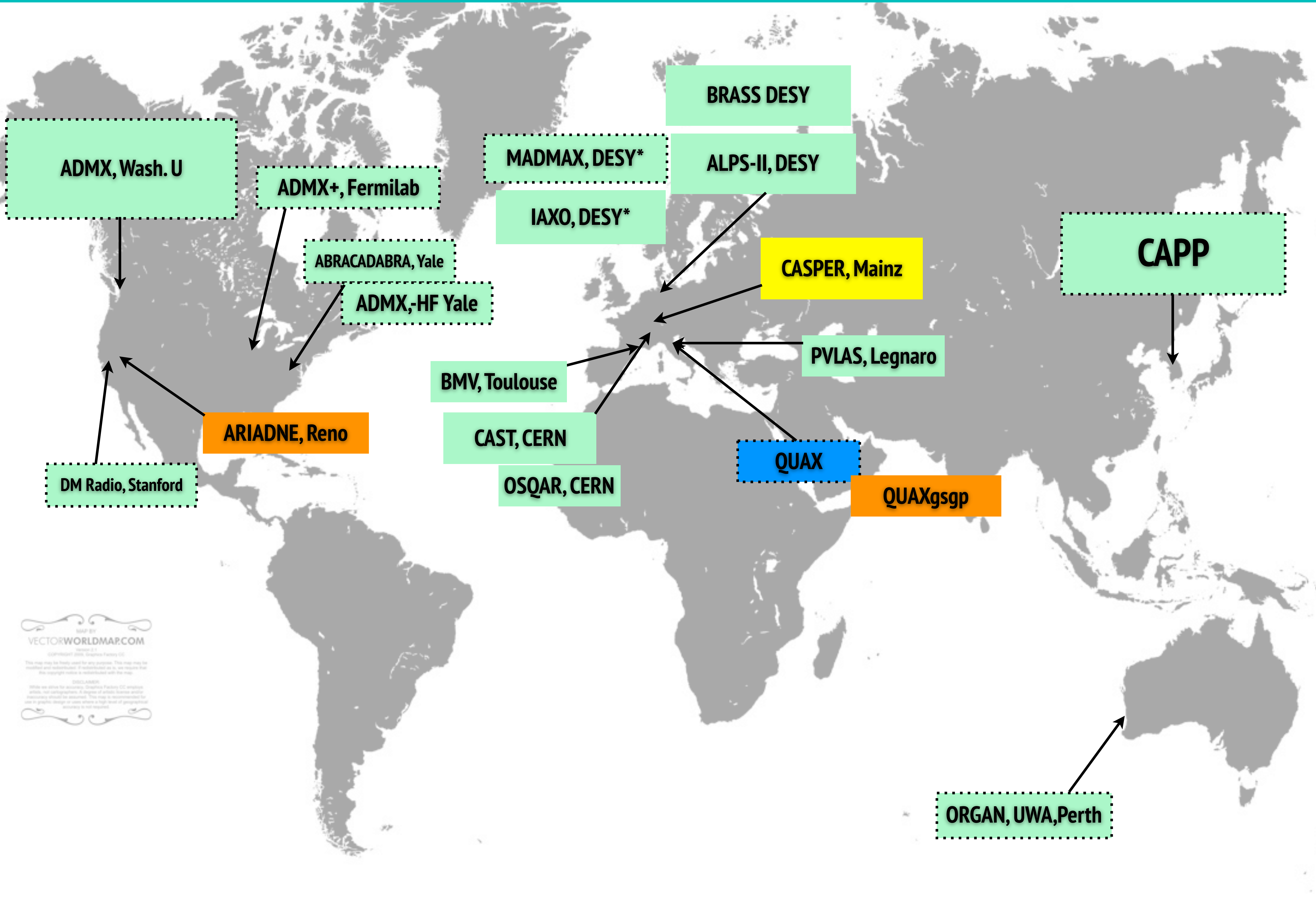




Lab experiments 2011

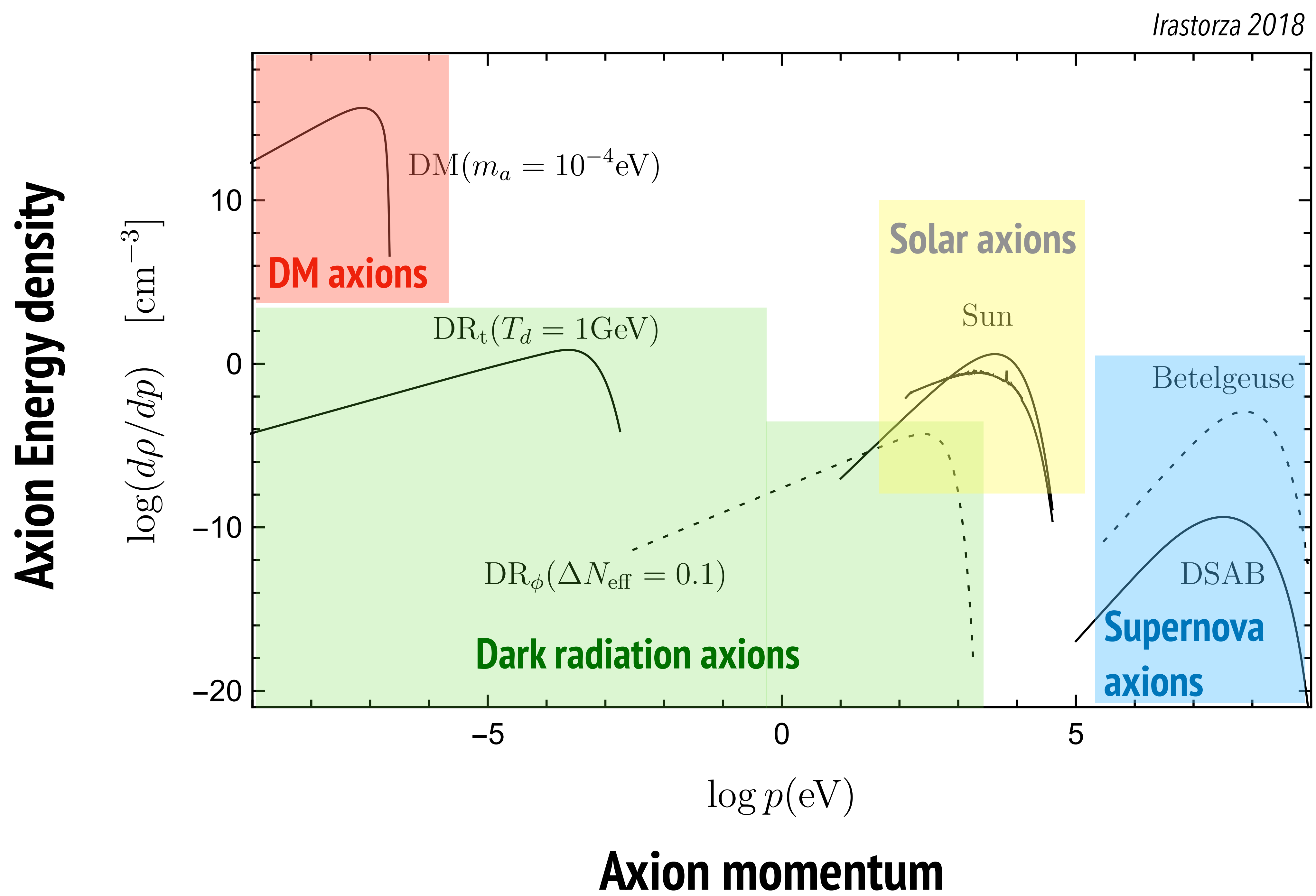


including R&D ~2017

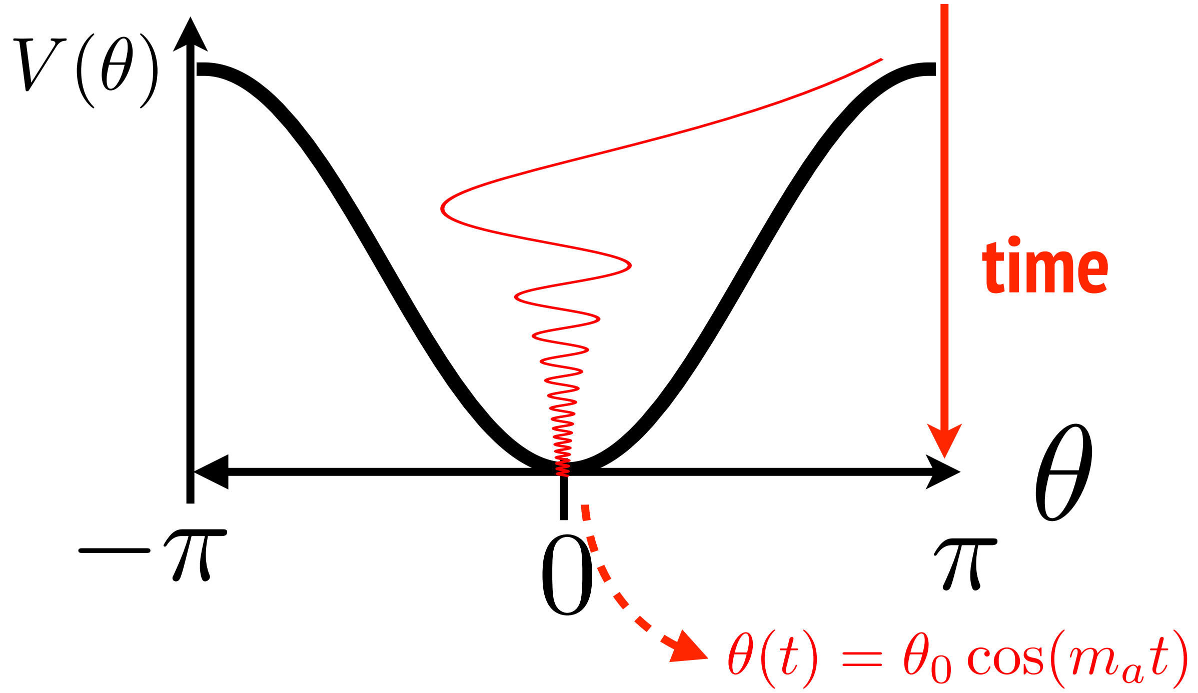


Search for Axions : Natural sources

- Naturally produced axions could be quite copious, save production and focus on detection!



Axion DM in the lab



Local Dark Matter density*

$$\rho_{\mathrm{aDM}} = 0.3 \frac{\mathrm{GeV}}{\mathrm{cm}^3}$$

$$\theta_0 = 3.6 \times 10^{-19}$$

Detecting Axion Dark Matter

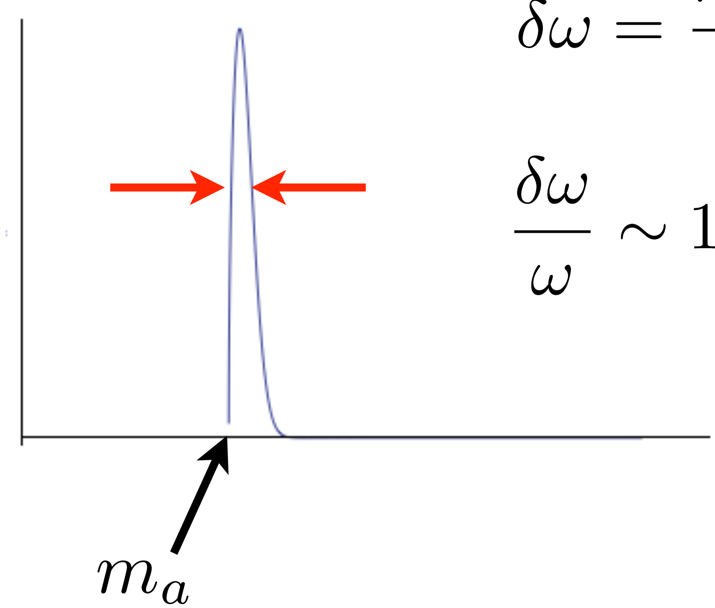
- $\theta_0 = 3.6 \times 10^{-19}$ **is a very small number but, oscillations allow for coherent detection!**

- **Axion spectrum is not exactly monochromatic, non-zero velocity of DM in the galaxy -> finite width**

frequency $\omega \simeq m_a(1 + v^2/2 + ...)$

$$\delta\omega = \frac{m_a v^2}{2}$$

$$\frac{\delta\omega}{\omega} \sim 10^{-6}$$



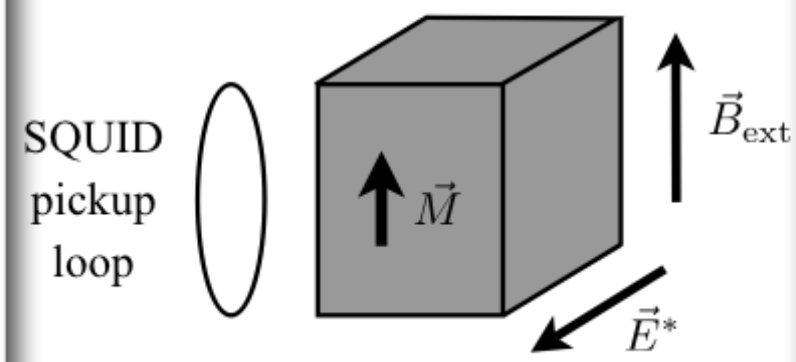
coherence time

$$\delta t \sim \frac{1}{\delta\omega} \sim 0.13\text{ms}\left(\frac{10^{-5}\text{eV}}{m_a}\right)$$

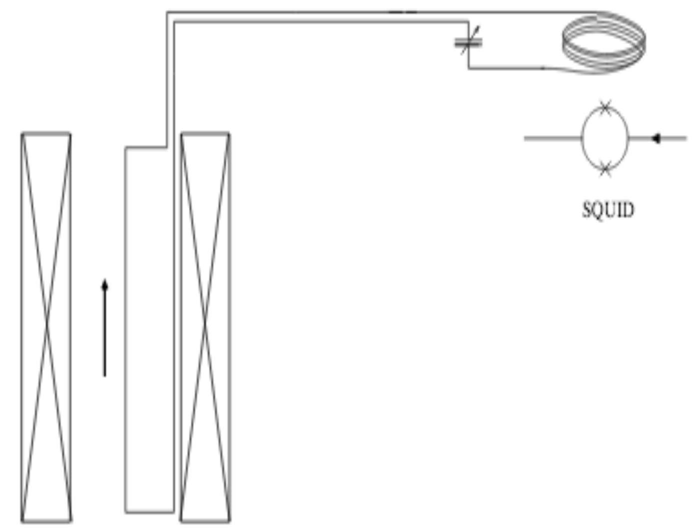
coherence length

$$\delta L \sim \frac{1}{\delta p} \sim 20\text{m}\left(\frac{10^{-5}\text{eV}}{m_a}\right)$$

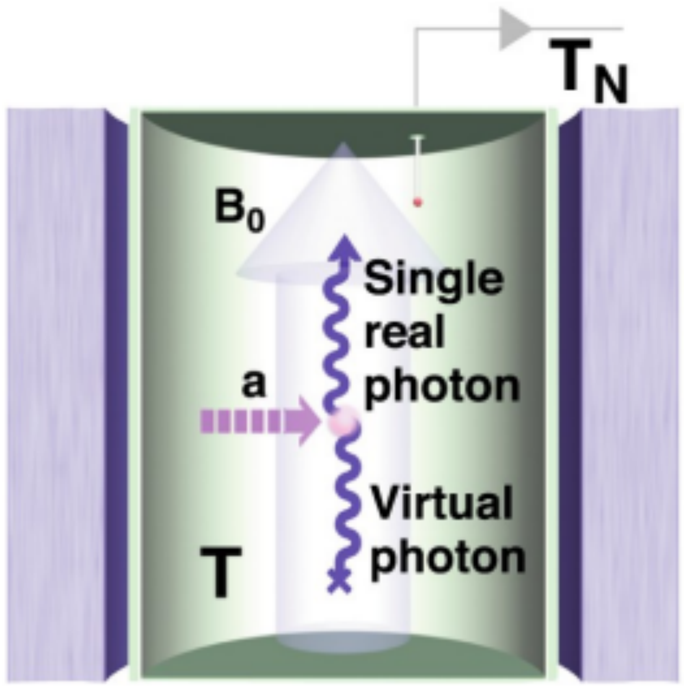
Oscillating EDM



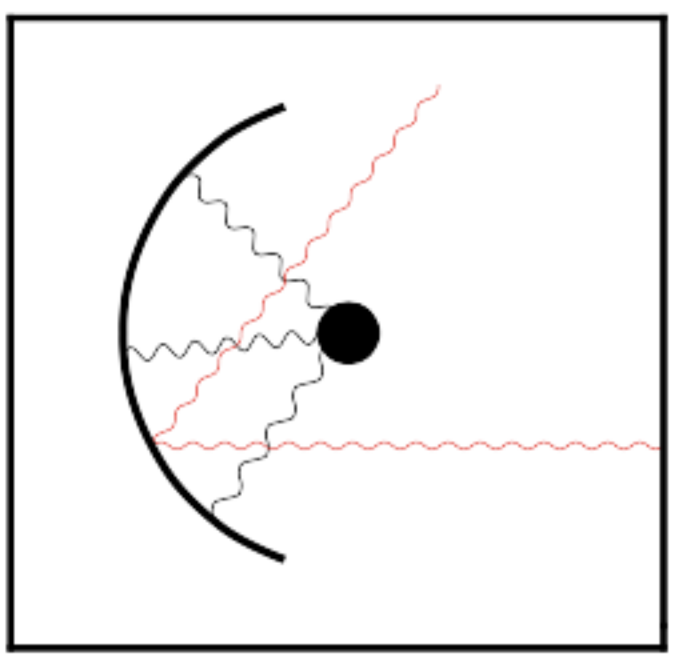
LC-circuit



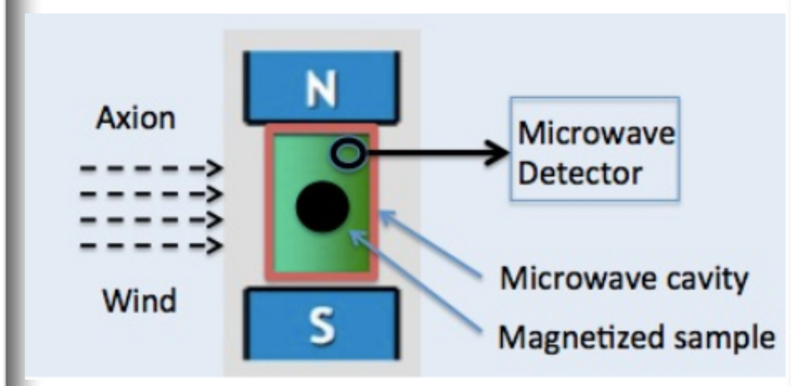
Cavities



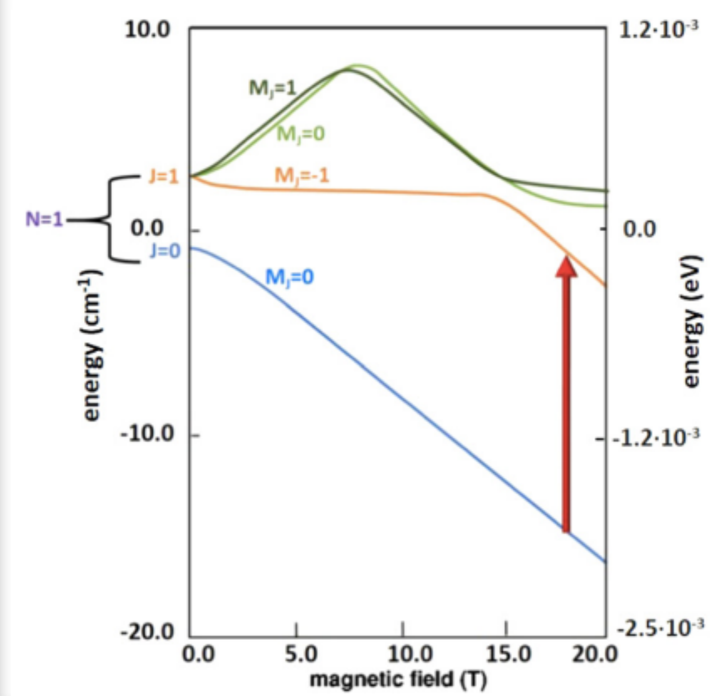
Mirrors



Ferromagnetic resonance



Atomic transitions



CASPER : oscillating EDM with NMR

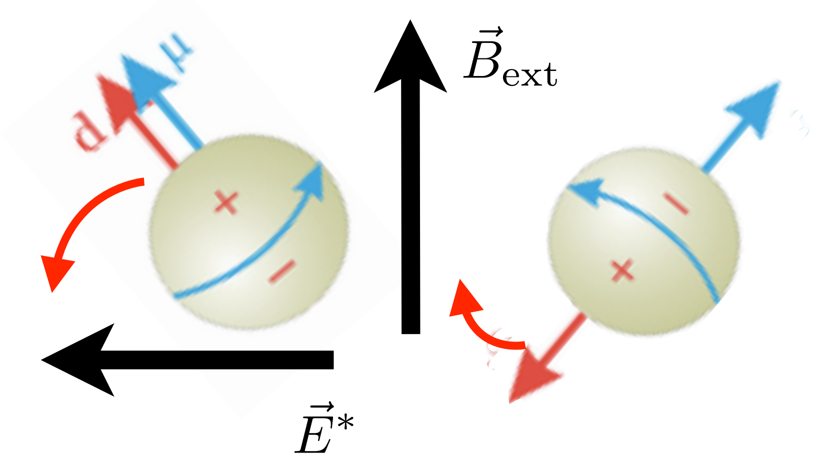
Mainz, Berkeley

D. Budker

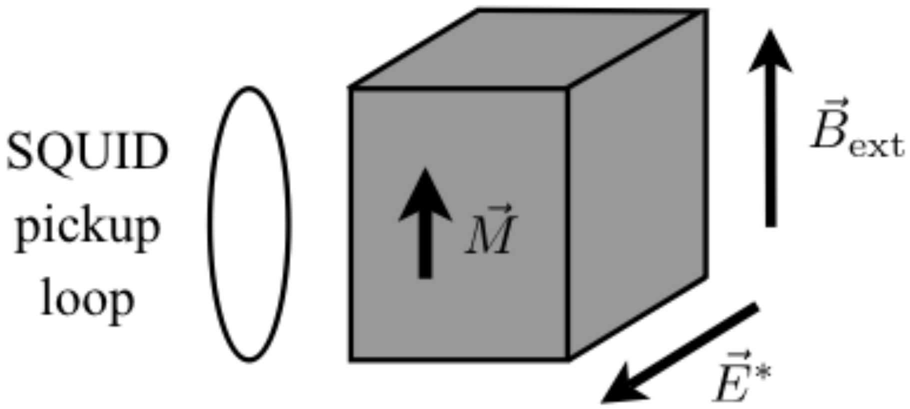
S. Rajendran

P. Graham

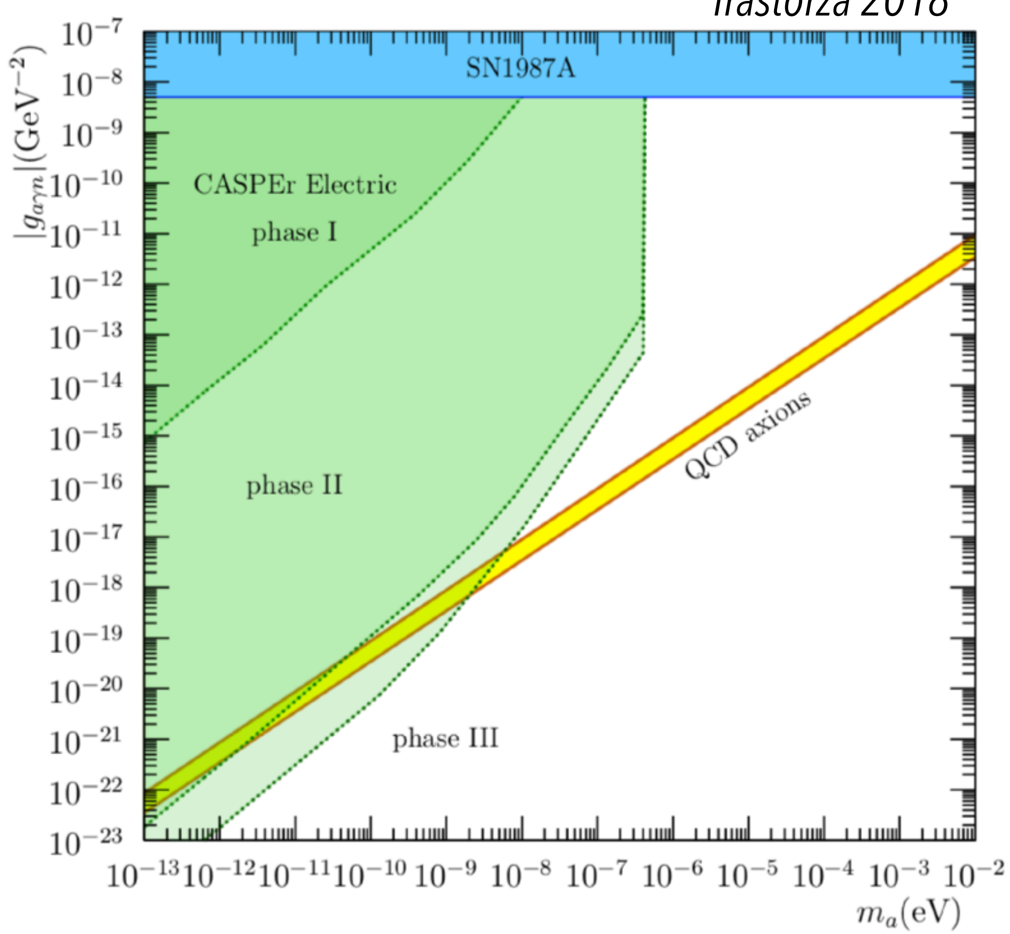
- **Oscillating neutron EDM** $d_n = -4 \times 10^{-3} \times \theta_0 \cos(m_a t)\ [\mathrm{e\,fm}]$



Oscillating EDM, effects add up, transverse magnetisation grows on resonance $m_a = \omega = \mu|\vec{B}_{\rm ext}|$



Irastorza 2018



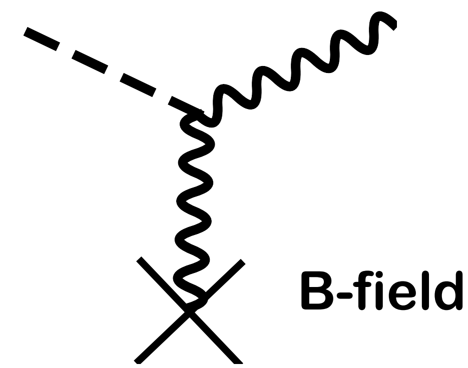
- *EDM + Large E-fields in PbTiO3*
- *Scan over frequencies, with Bext*
- *Mainz (D. Budker's group) & Berkeley*
- *Phase I starts in 2017, Phase II physics results …*
- *Mass range limited by B-field strength*

Axion DM in a B-field

- Axion photon coupling in a strong B-field becomes a source of E-field

$$\mathcal{L}_I = -C_{a\gamma}\frac{\alpha}{2\pi}\theta(t)\,\mathbf{B}_{\rm ext}\cdot\mathbf{E}$$

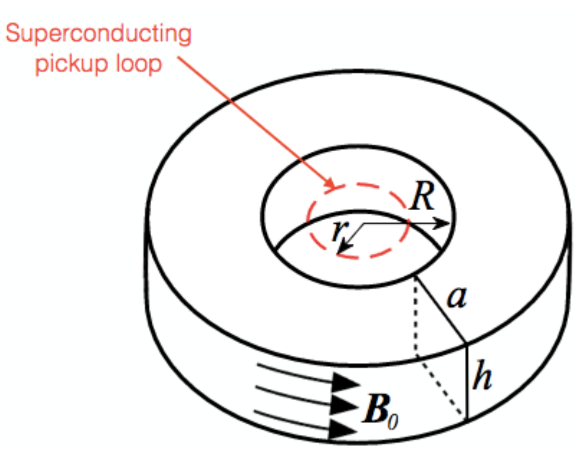
source



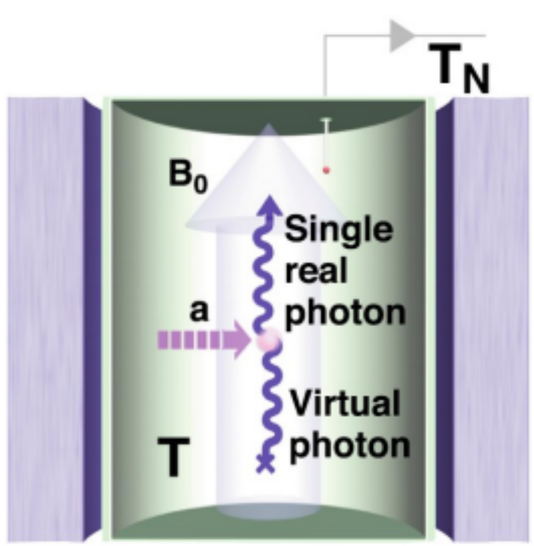
E-field $E \sim \mathcal{O}(10^{-12}\mathrm{V/m})\frac{|\mathrm{B}_{\rm ext}|}{10\,\mathrm{T}}C_{a\gamma}\times\cos(m_a t)$

Power $P/Area \sim |\mathbf{E}_a|^2 \sim 2\times 10^{-27}\left(\frac{\mathrm{B}}{5\mathrm{T}}\frac{C_{a\gamma}}{2}\right)^2\frac{\mathrm{Watt}}{1\,\mathrm{m}^2}$

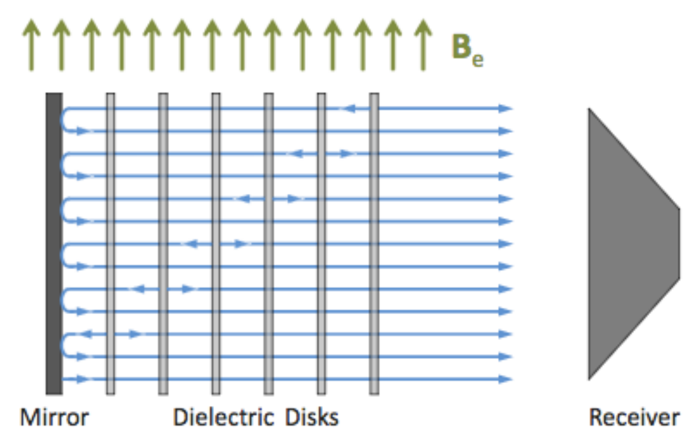
- Four different techniques:



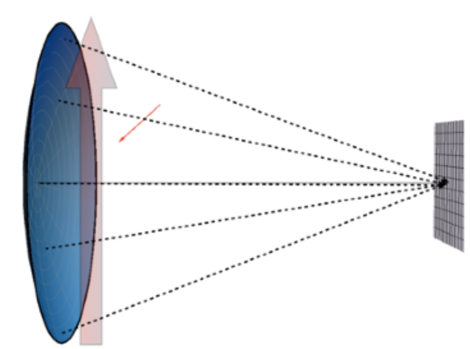
DM Radio



Cavities



Dielectric haloscope

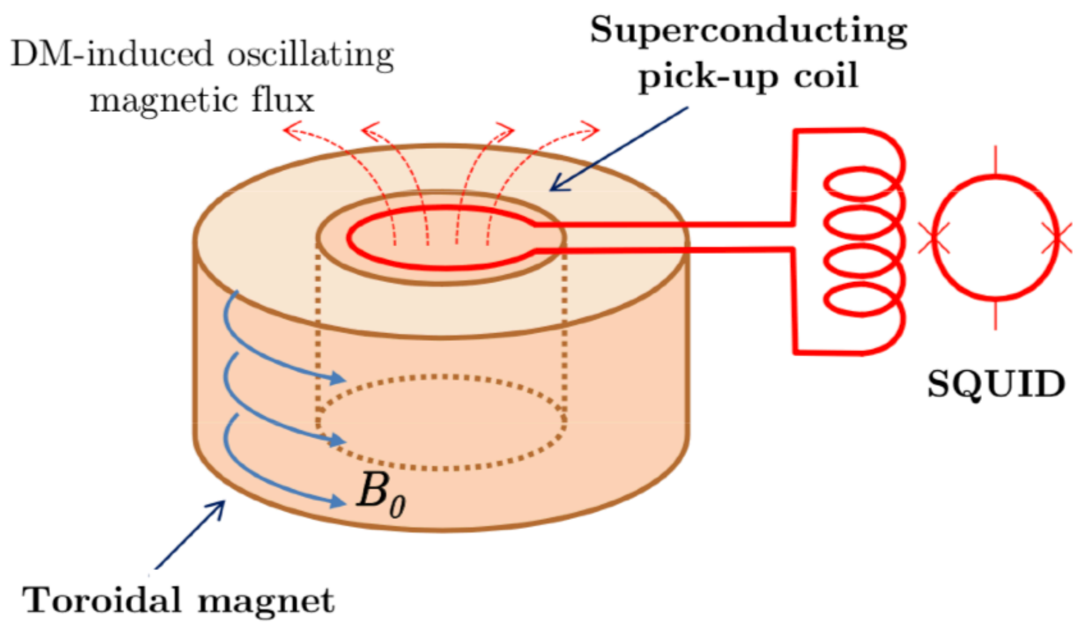


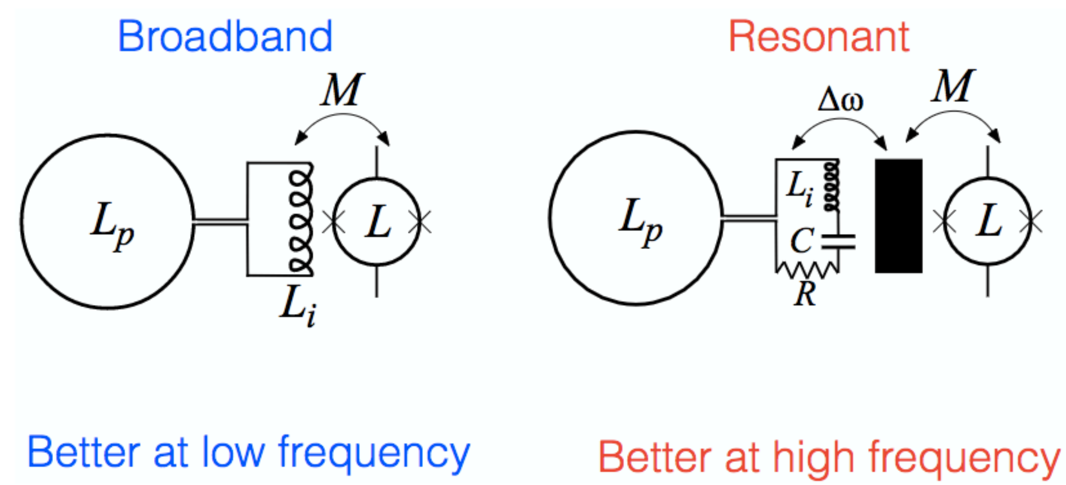
Dish antenna

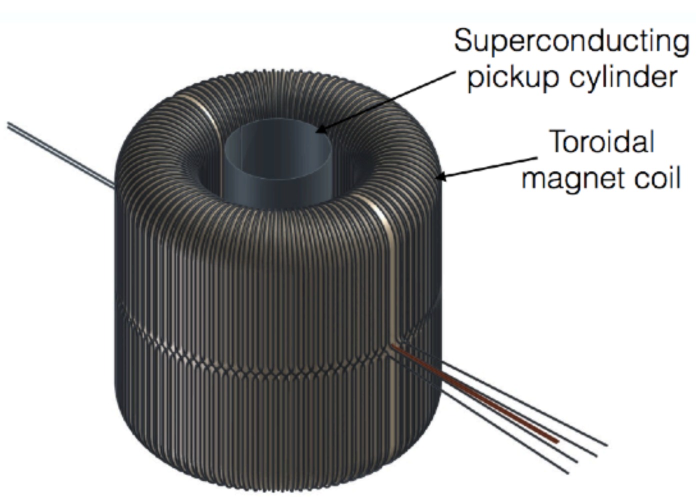
DM Radio

- **Toroidal axion-induced E-field generates oscillating B-field along z**

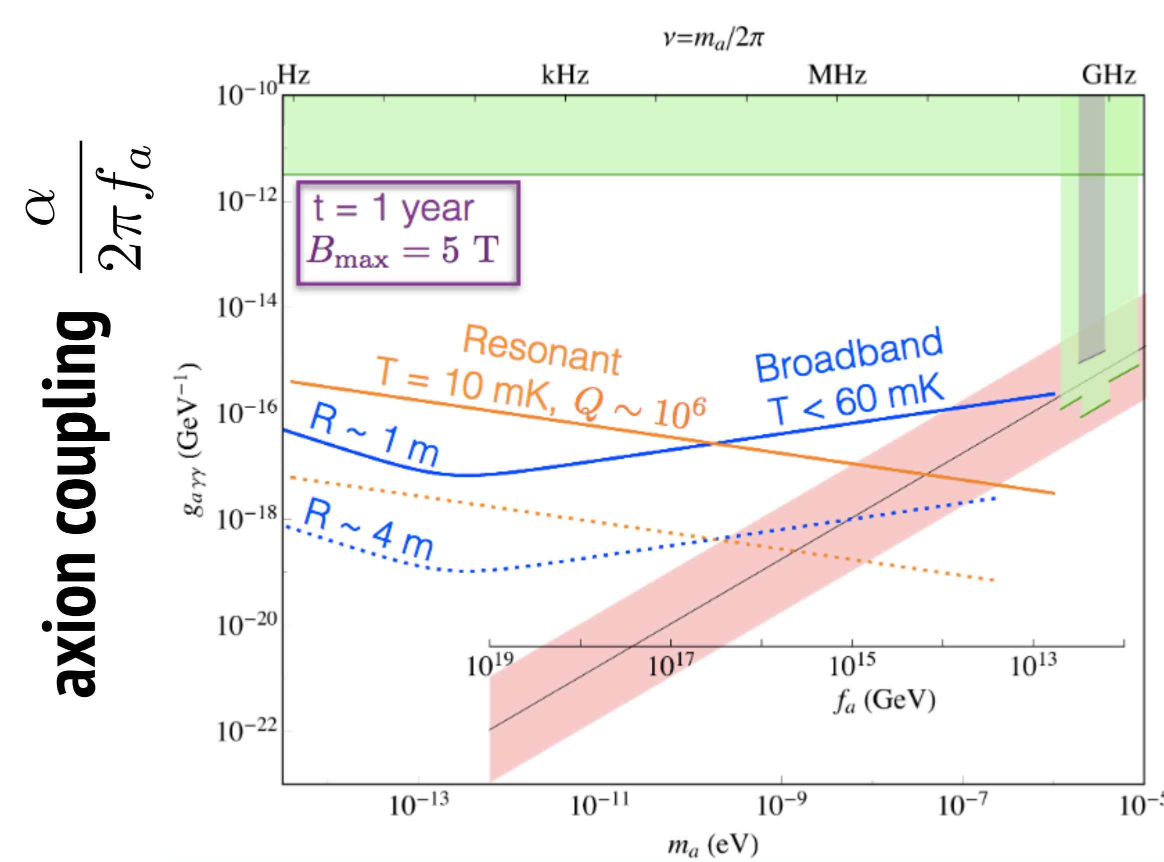
Sikivie PRL 112 (2014)
Chaudhuri PRD92 (2015)
Kahn PRL 117 (2016)







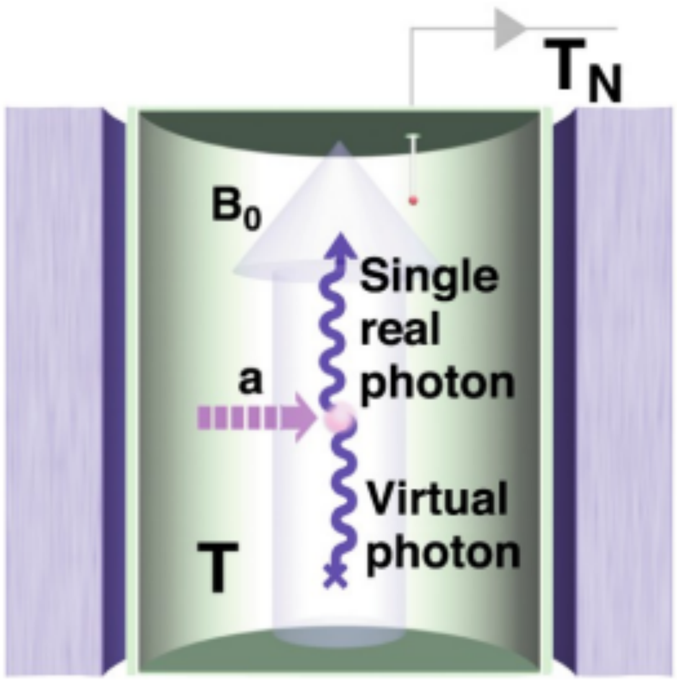
ABRACADABRA (MIT)
10 cm, 1m , 4m ...



Resonant cavities: haloscopes

P. Sikivie

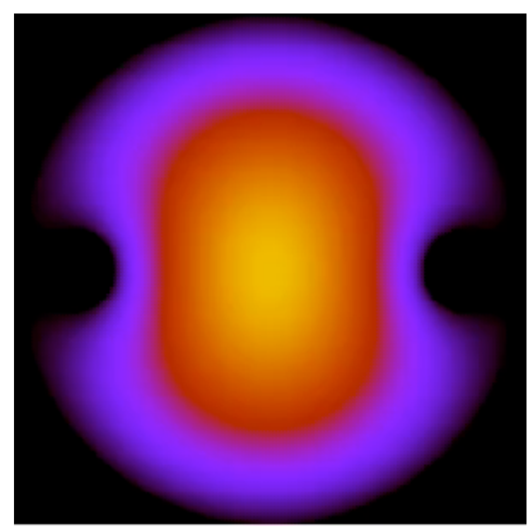
- Boost the axion-generated E-field in a tuned resonant cavity

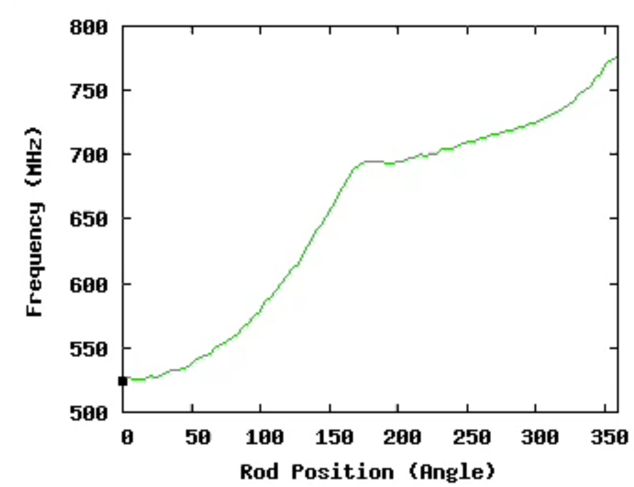


$$P_{\rm out} \sim Q|\mathbf{E}_a|^2 V m_a$$

- Cavity quality factor $Q \sim 10^5$
- B-fields $B \sim 10\mathrm{T}$
- Volume $\sim 1/m_a^3$ (typically a few liters)
- Temperature $T \sim 0.2 - 4\,\mathrm{K}$
- System T ~ Quantum limited (SQUID, JPA)

Scanning over frequencies





- At high freq. limited by small volume and high noise
- At low freq. by getting a large enough B-field

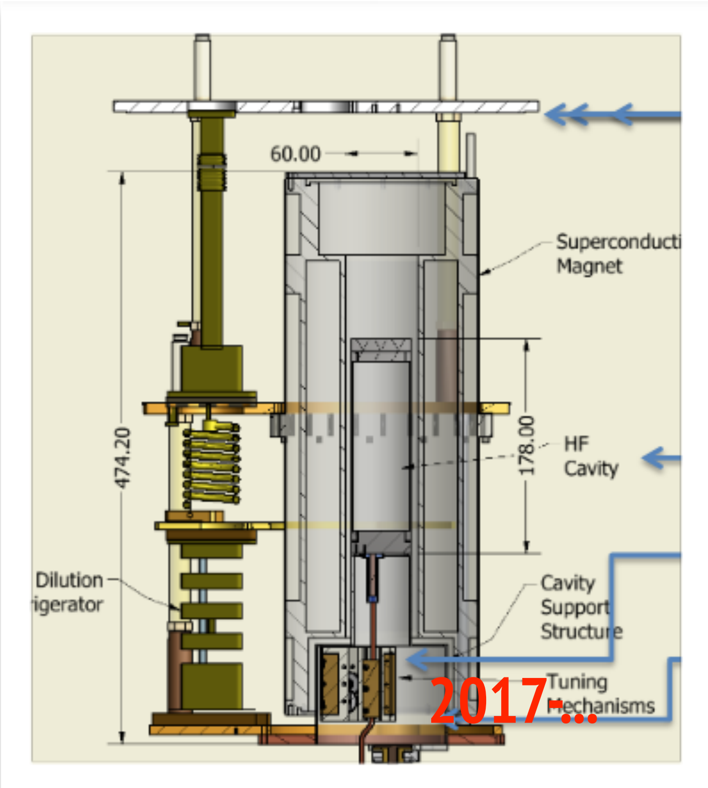
Cavity experiments

- Physical dimensions $L \sim 1/m_a$

ADMX-Seattle

new data!!!

CULTASK - CAPP -Korea



2017-...

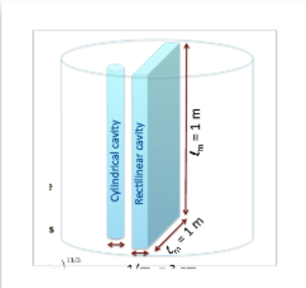
ORGAN-UWA Perth

2017-...

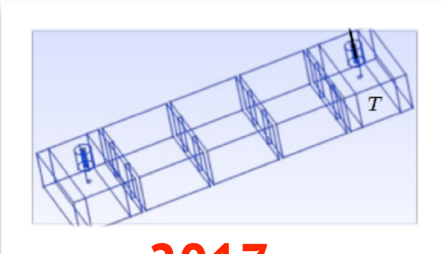
HAYSTAC-Yale

2016-...

ADMX-Fermilab



RADES



2017-...

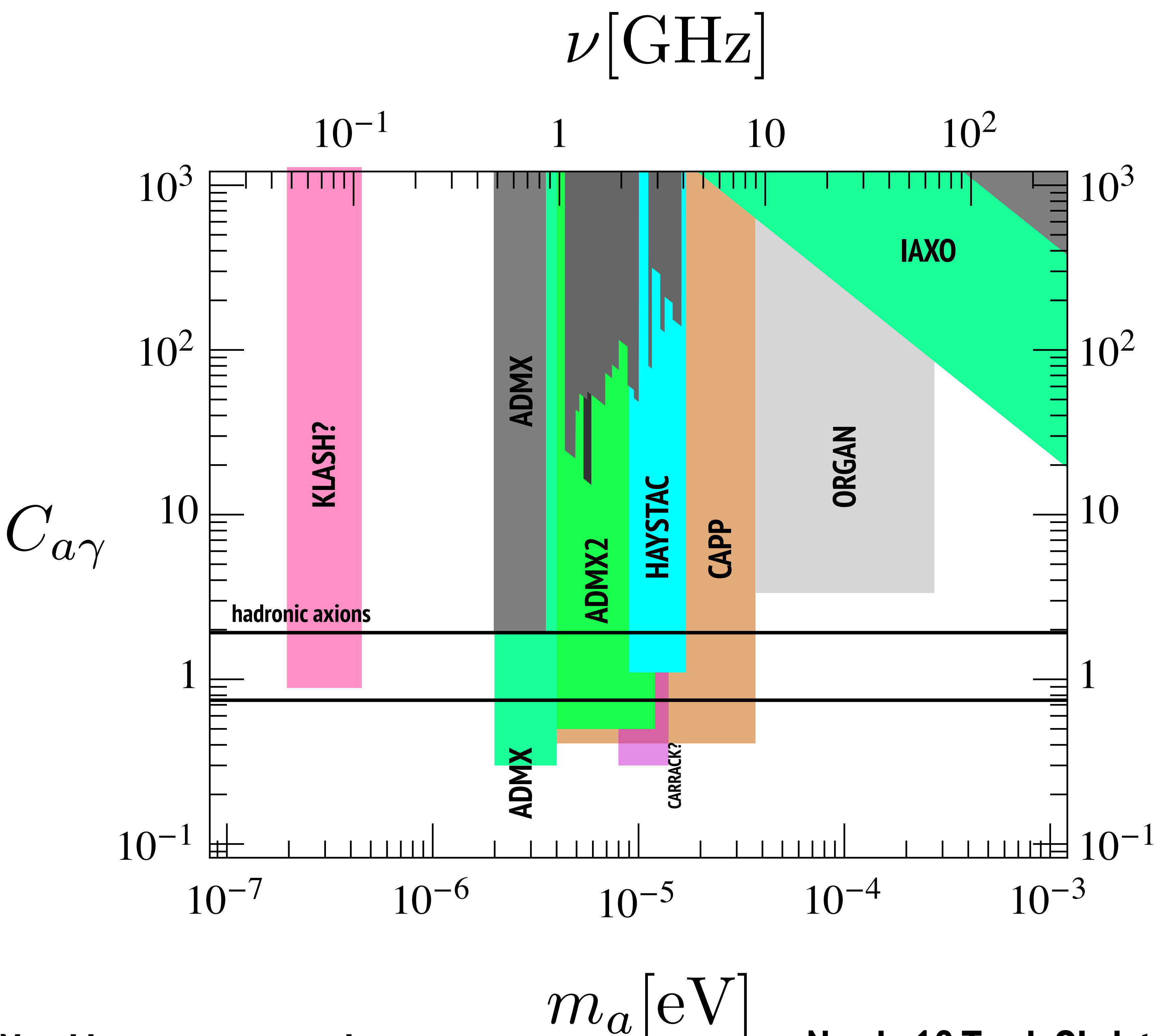
CAST-CAPP

2017-...

KLASH?

??-...

Projected optimistic sensitivities



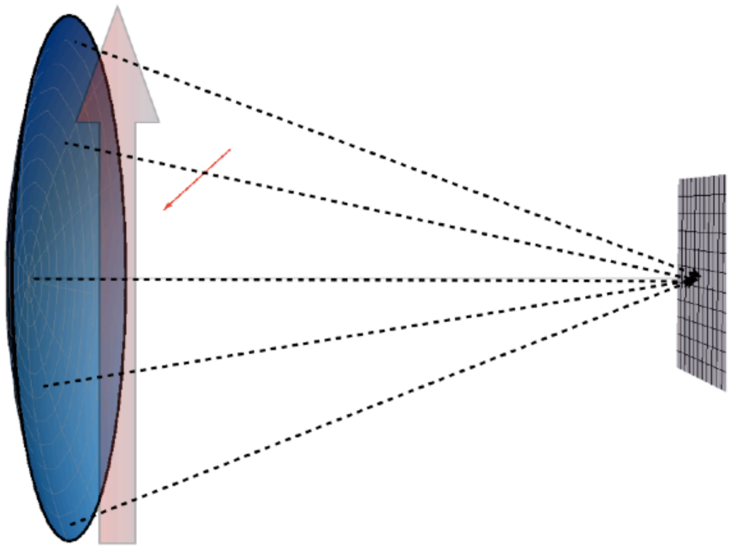
- Need larger magnet volume

- Need >10 T, sub QL detection, Q~10^6

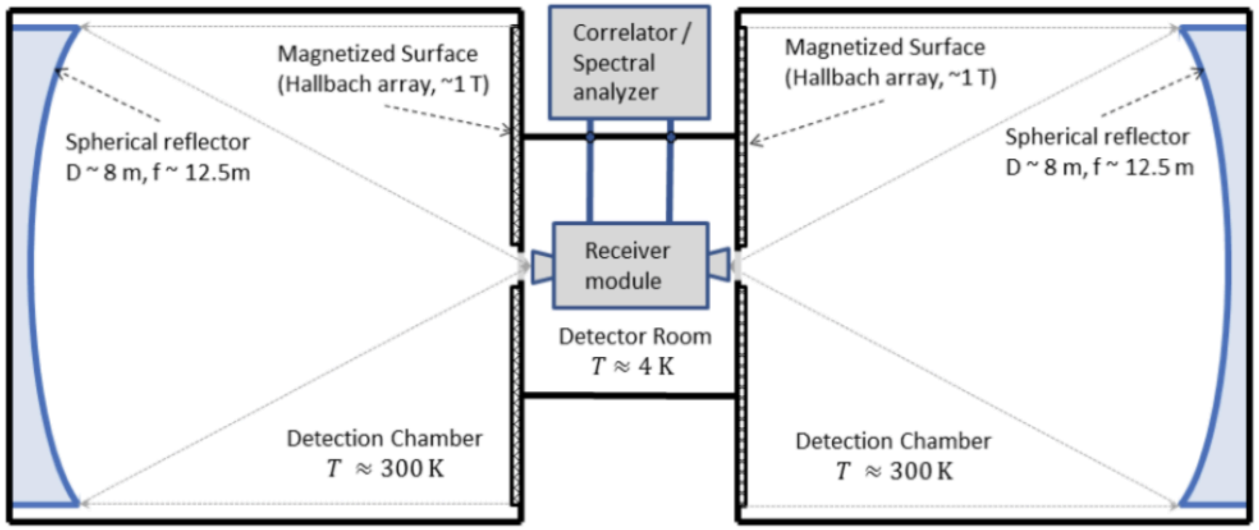
- or combine many cavities ...

Dish antenna

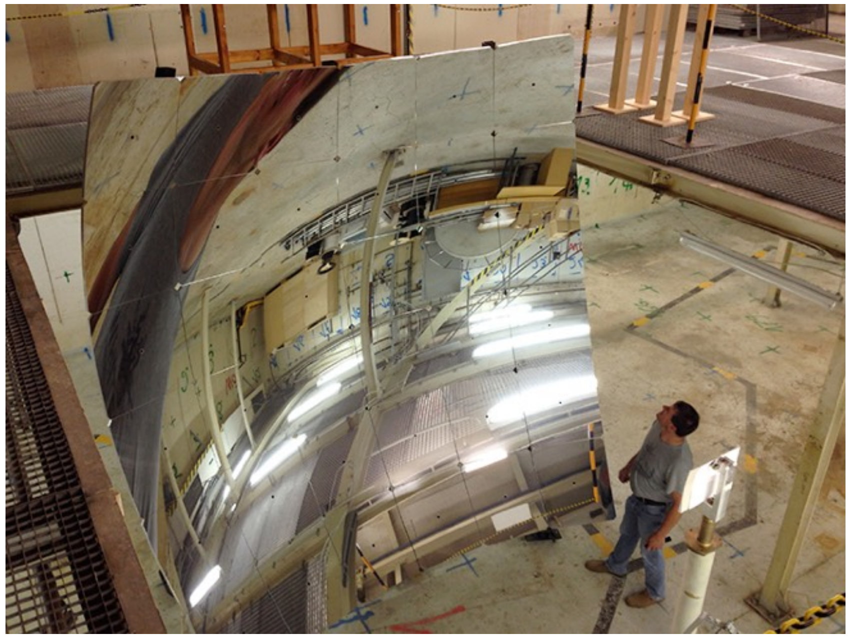
- Detect radiated power from a huge ($Am_a^2 \gg 10^6$) magnetised dish
- Broadband, no resonance enhancement; Only detector needs to be at T~mK (high reflectivity dish)
- Magnetise Area with permanent-magnets, photon counting?



$$P/Area \sim |\mathbf{E}_a|^2 \sim 2 \times 10^{-27} \left(\frac{\mathrm{B}}{5\mathrm{T}} \frac{C_{a\gamma}}{2}\right)^2 \frac{\mathrm{Watt}}{1\,\mathrm{m}^2}$$



BRASS @ Hamburg



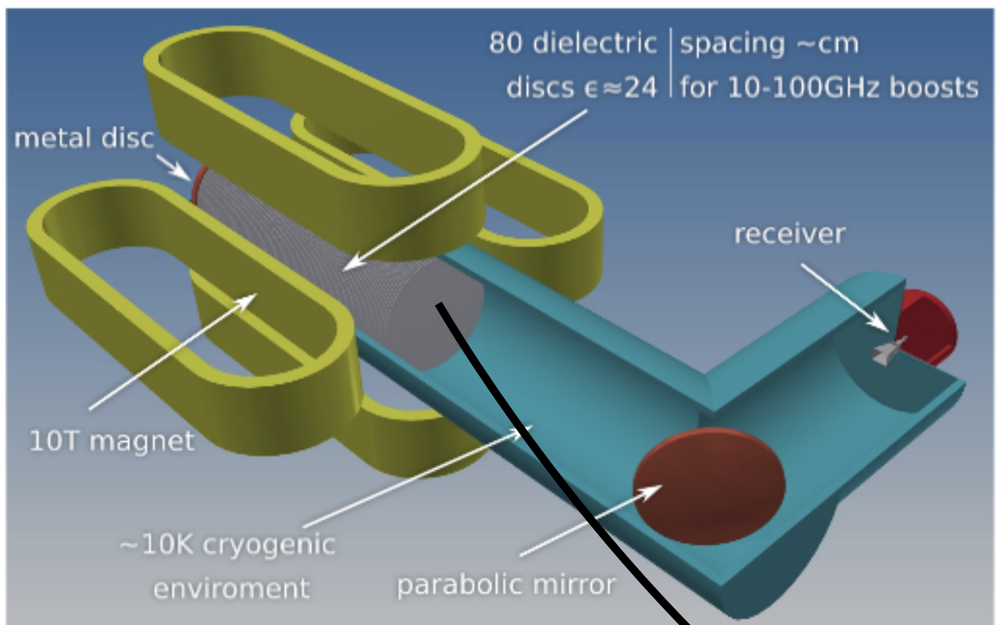
FUNK experiment (KIT)

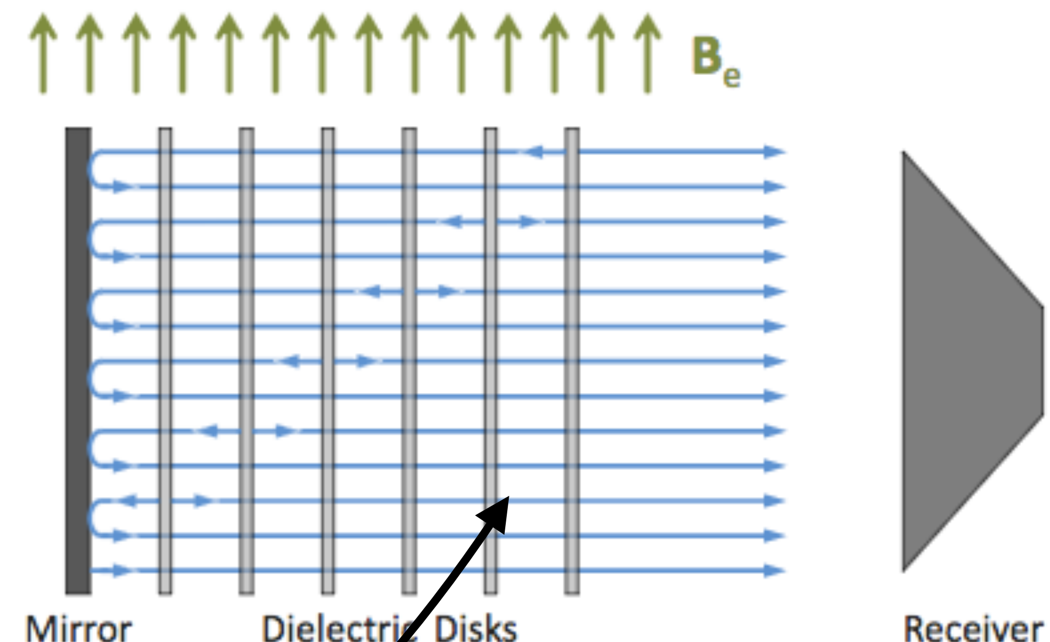
Dielectric haloscope : MADMAX

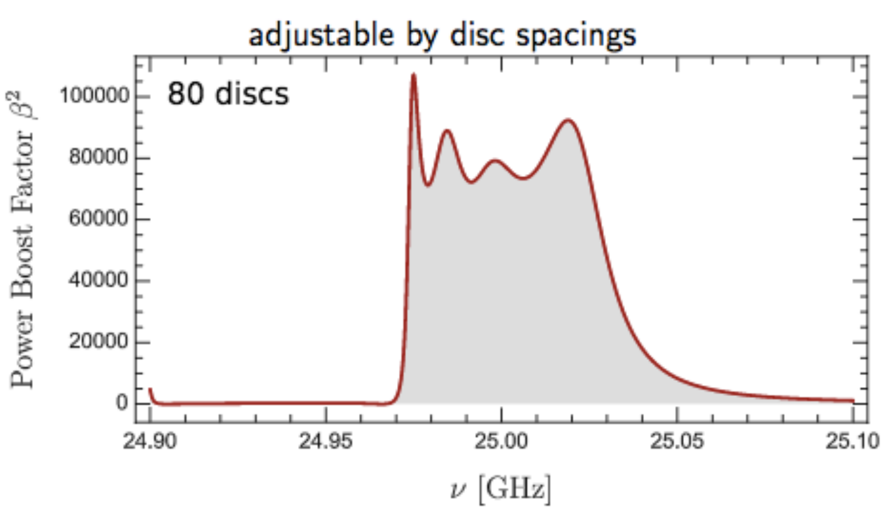
- Hybrid system, large area + multiple emitters + a bit of resonant enhancement

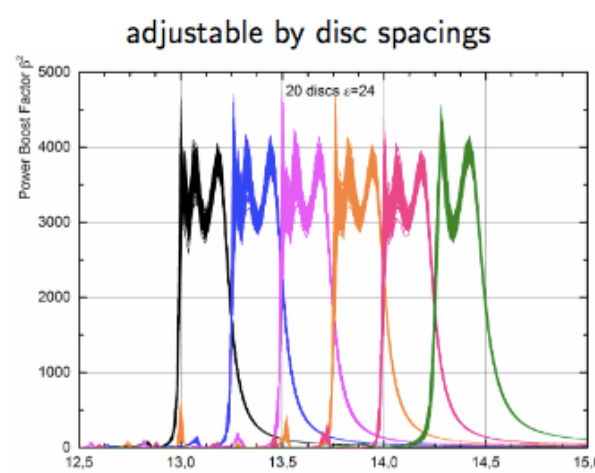
$$\frac{P}{Area} \sim 2\times 10^{-27}\,\frac{\mathrm{W}}{\mathrm{m}^2}\left(\frac{c_\gamma}{2}\,\frac{B_{||}}{5\mathrm{T}}\right)^2\frac{1}{\epsilon}\times\beta(\omega)$$

boost factor



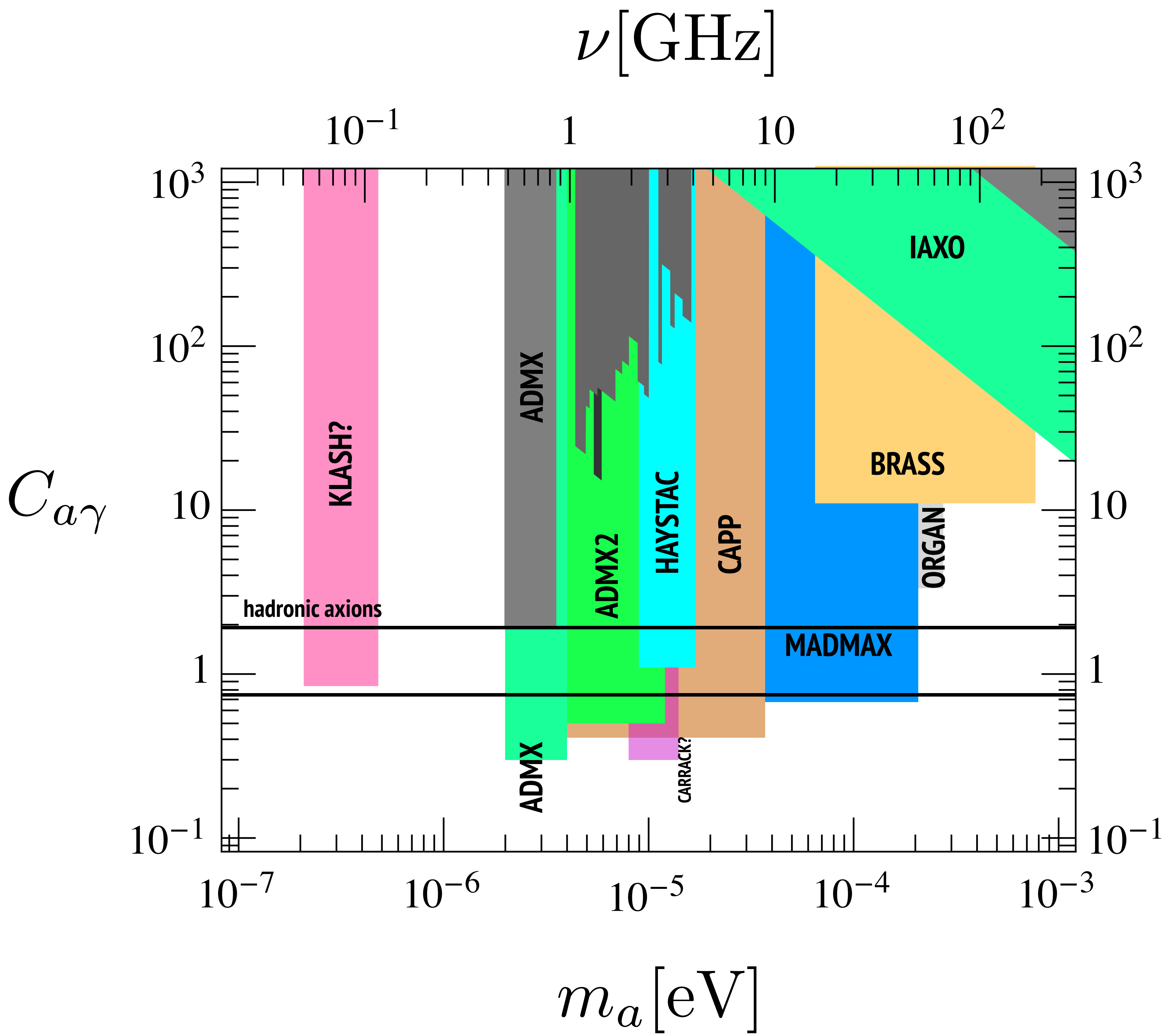




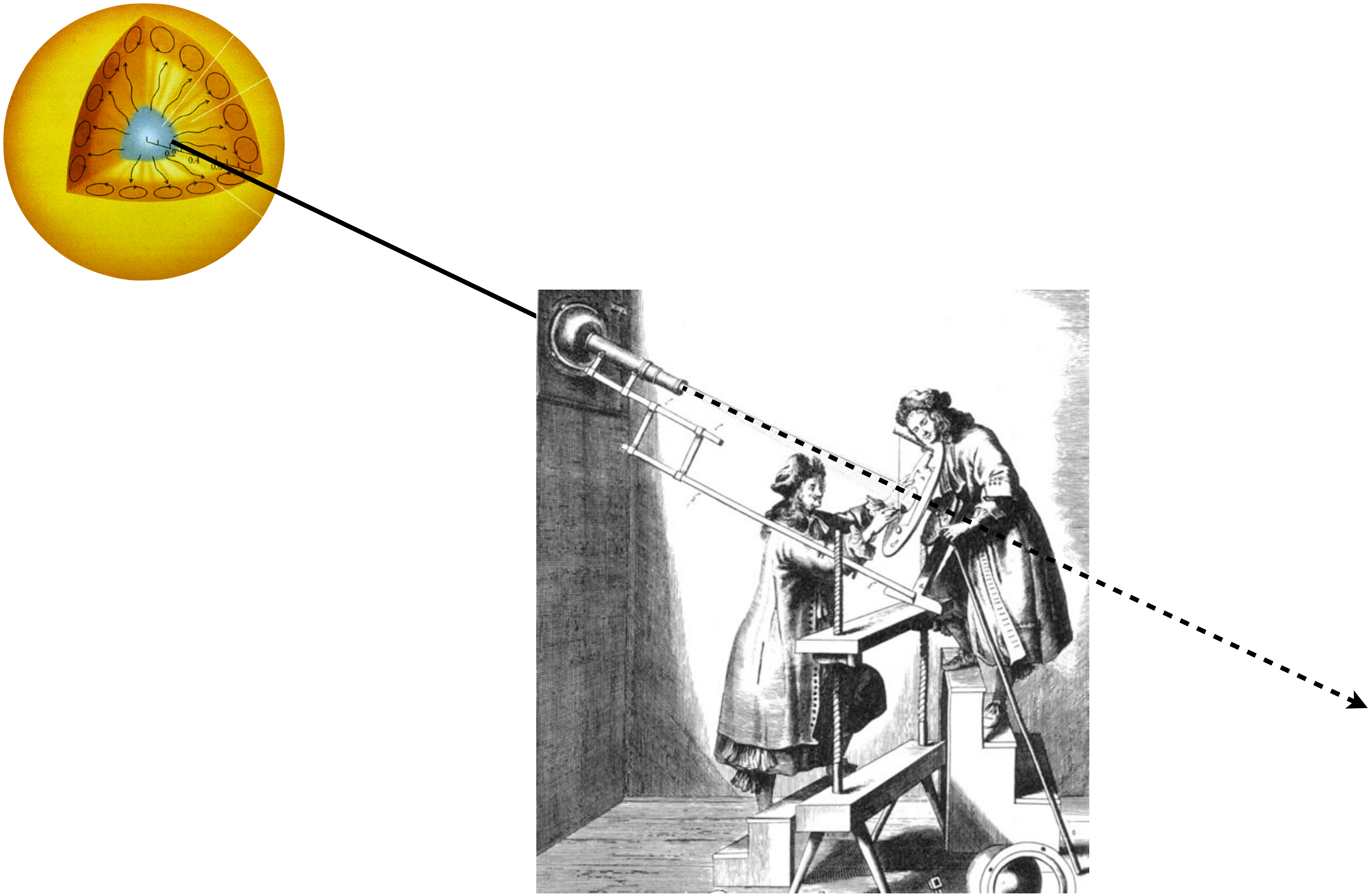


MADMAX: MAgnetised Disk and Mirror Axion eXperiment: MPP Munich, Hamburg Uni, DESY, Saclay, Zaragoza U

Projected sensitivities

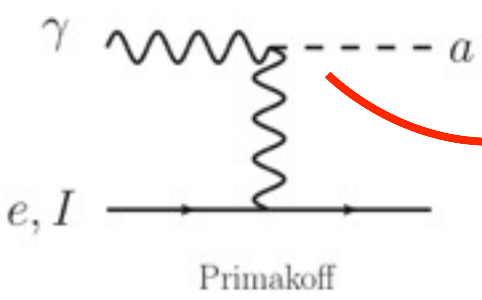


1.2 Detecting Solar Axions : Helioscopes



Axions from the Sun

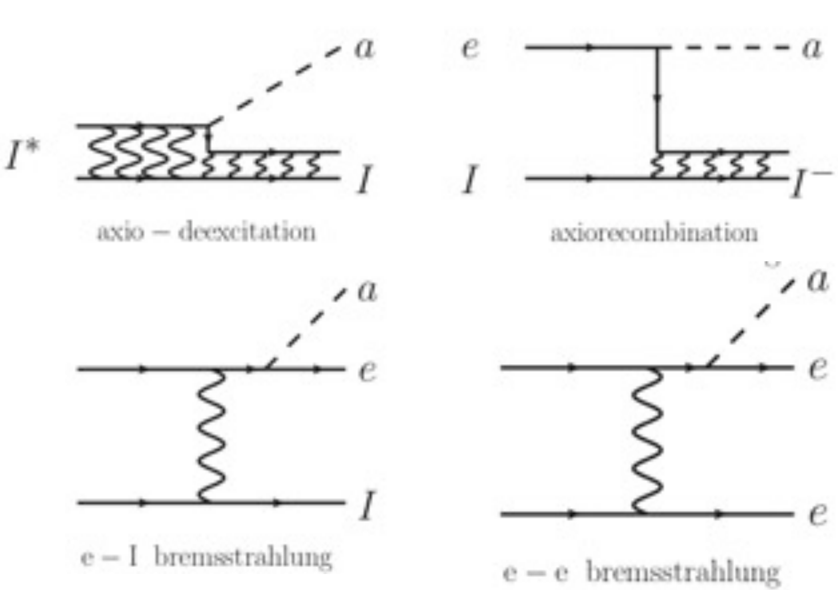
Hadronic axions (KSVZ)

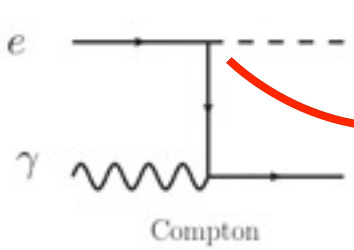


$$\frac{C_\gamma \alpha}{2\pi f_a} \frac{F_{\mu\nu}\widetilde{F}^{\mu\nu}}{4} a$$

$g_{a\gamma}$

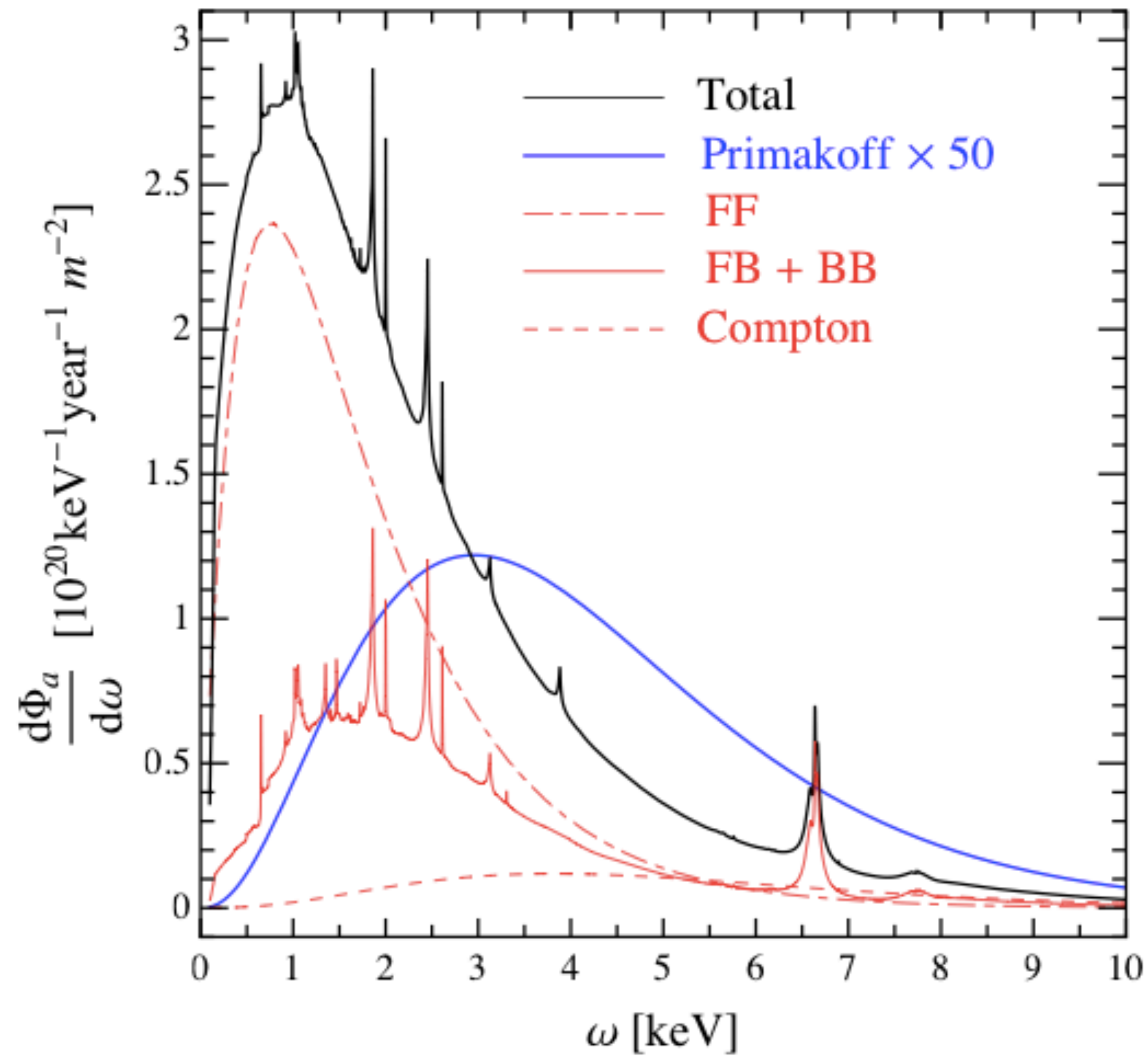
Non hadronic (DFSZ, e-coupling!)





$$\frac{C_e m_e}{f_a} i\bar{e}\gamma_5 e\, a$$

g_{ae}



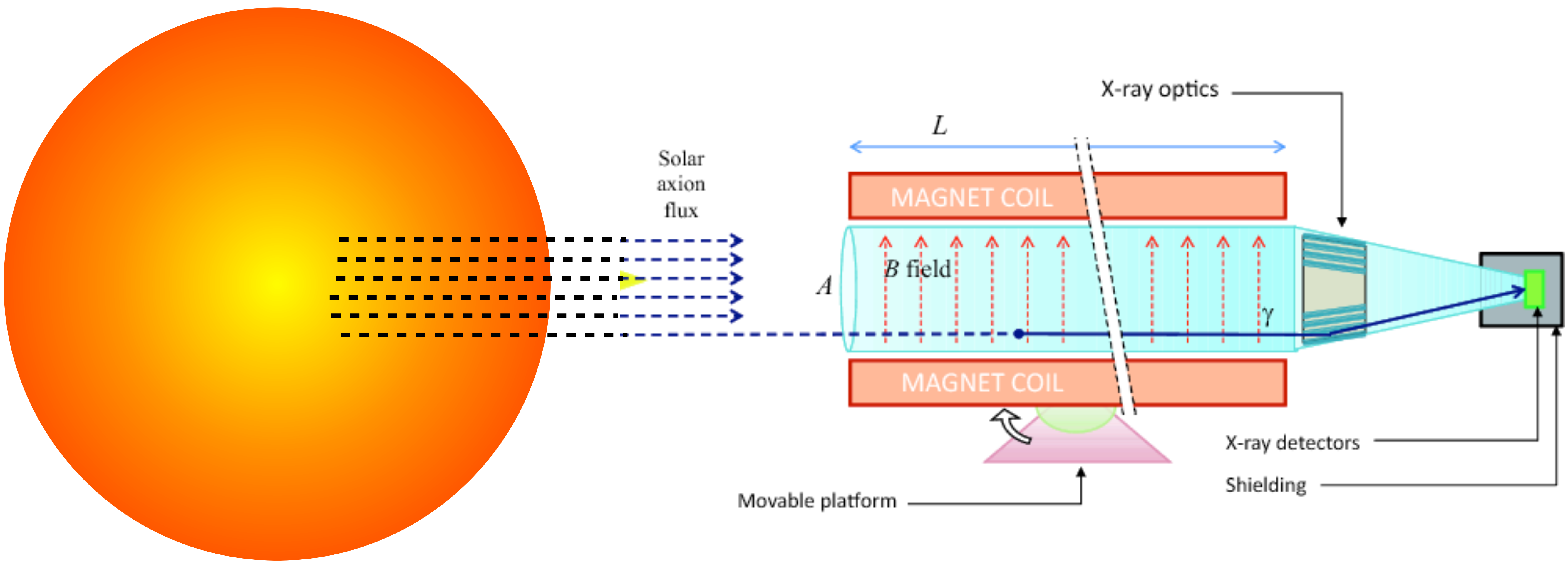
$g_{ae} = 10^{-13}$

$g_{a\gamma} = 10^{-12}\text{GeV}^{-1}$

typical of meV mass axions

Helioscopes

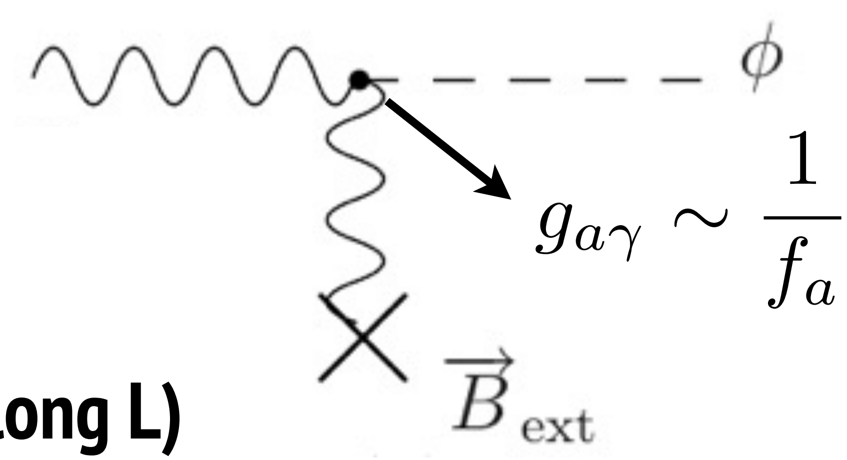
The Sun is a copious emitter of axions! **convert into X-rays** **focus** **detect**



Conversion probability

$$P(a \leftrightarrow \gamma) = \left(\frac{2 g_{a\gamma} B_T \omega}{m_a^2}\right)^2 \sin^2\left(\frac{m_a^2 L}{4\omega}\right)$$

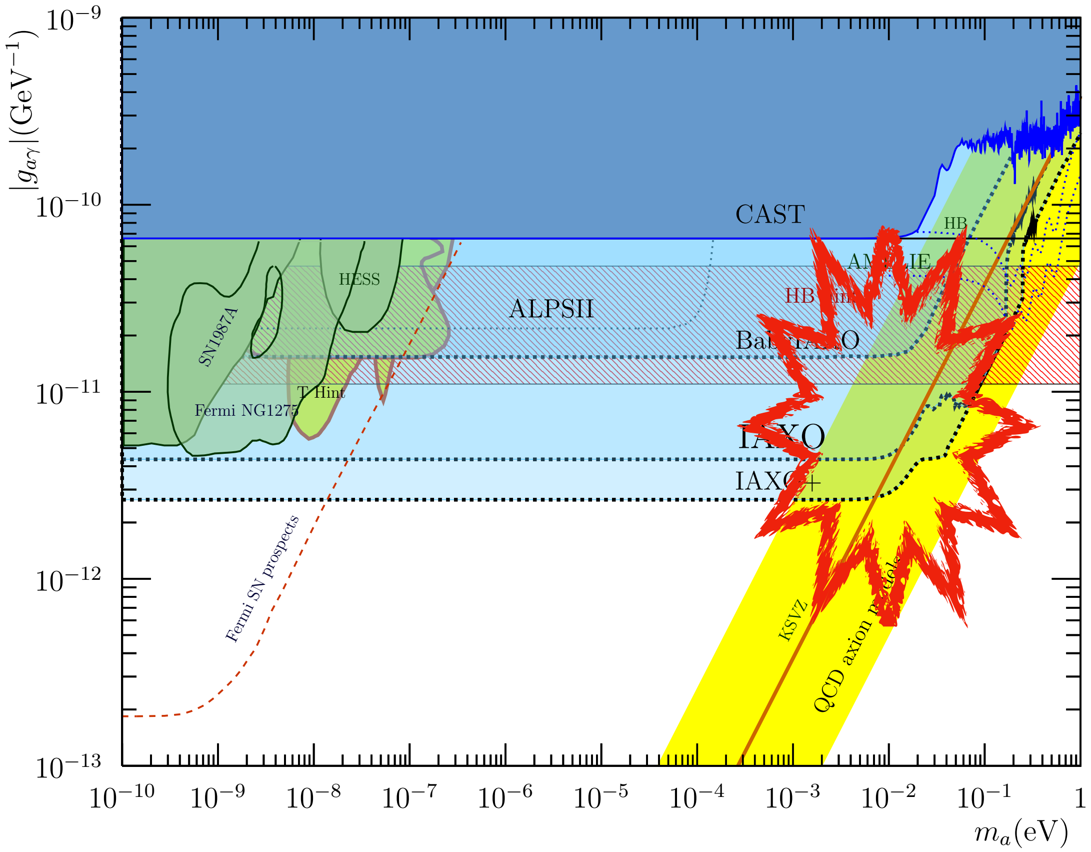
$$P(a \leftrightarrow \gamma) \sim 10^{-20}\left(\frac{B}{3\,\mathrm{T}}\frac{L}{20\,\mathrm{m}}\right)^2$$ **(coherence along L)**



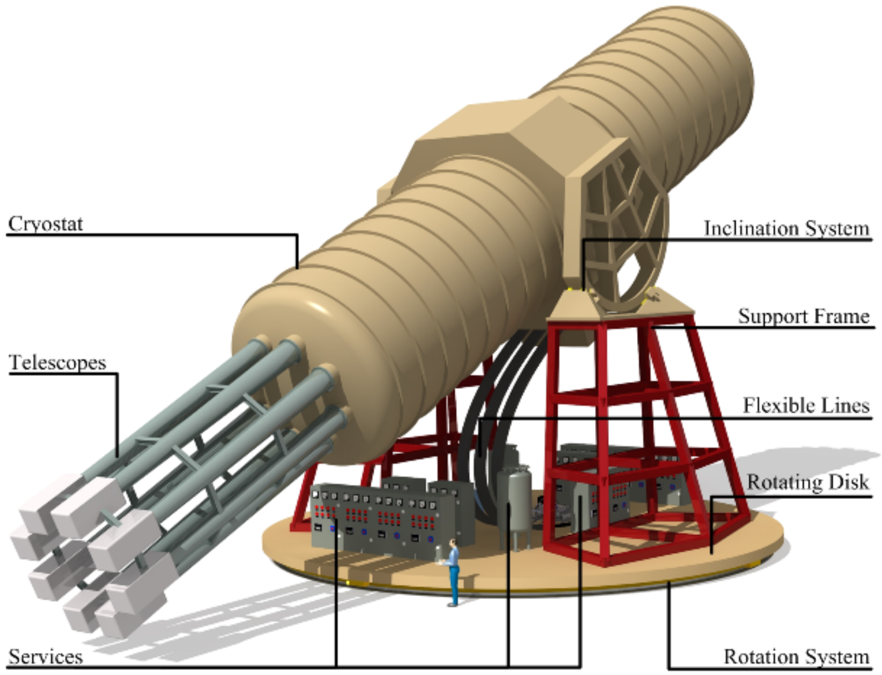
Past and the future

CAST (LHC dipole 9.3 m, 9T)

- 1~2 h tracking/day (sunset,dawn)
- 3 Detectors (2 bores)
- X-ray optics
- small aperture

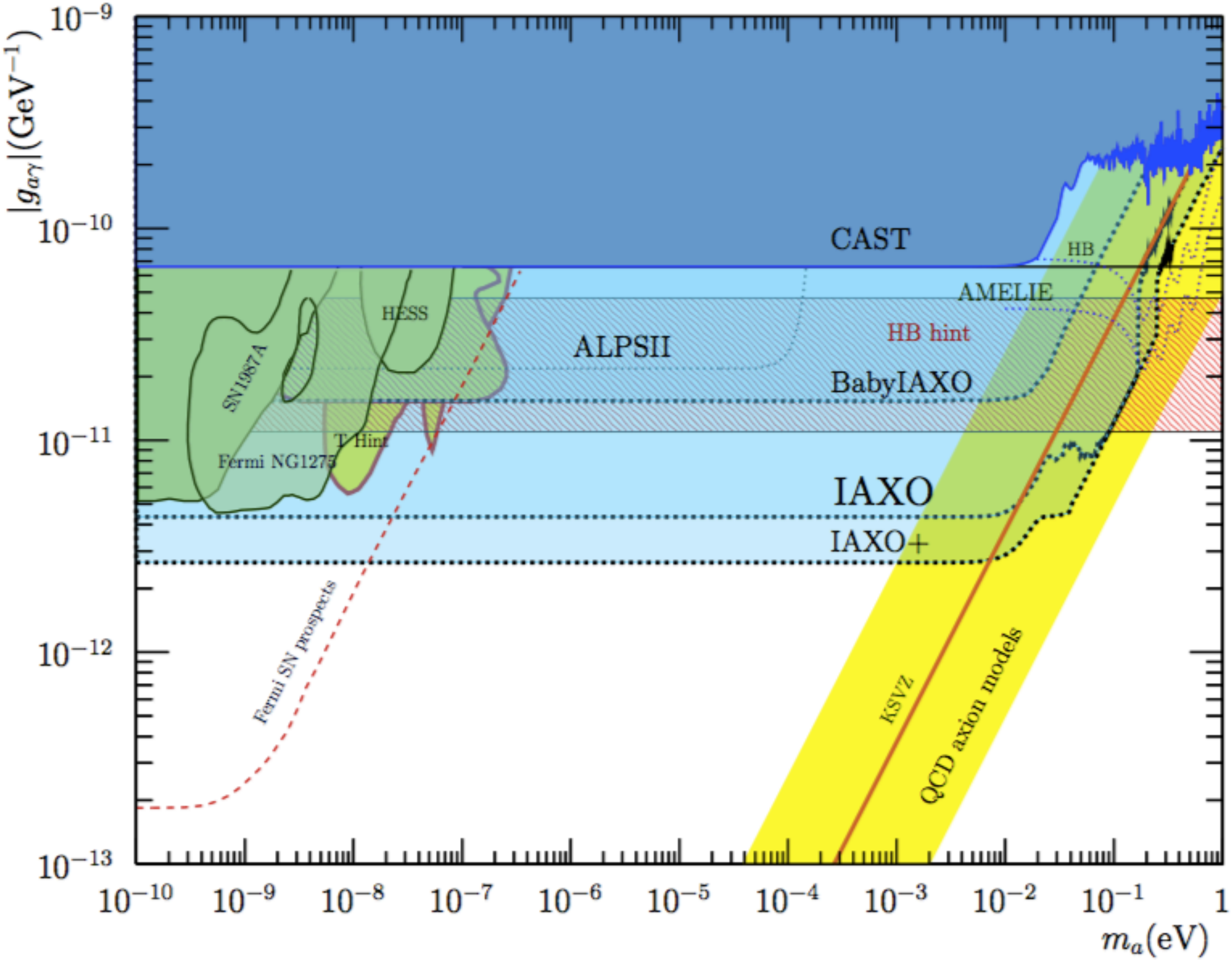


IAXO (proposed toroid) 20 m, 3T)

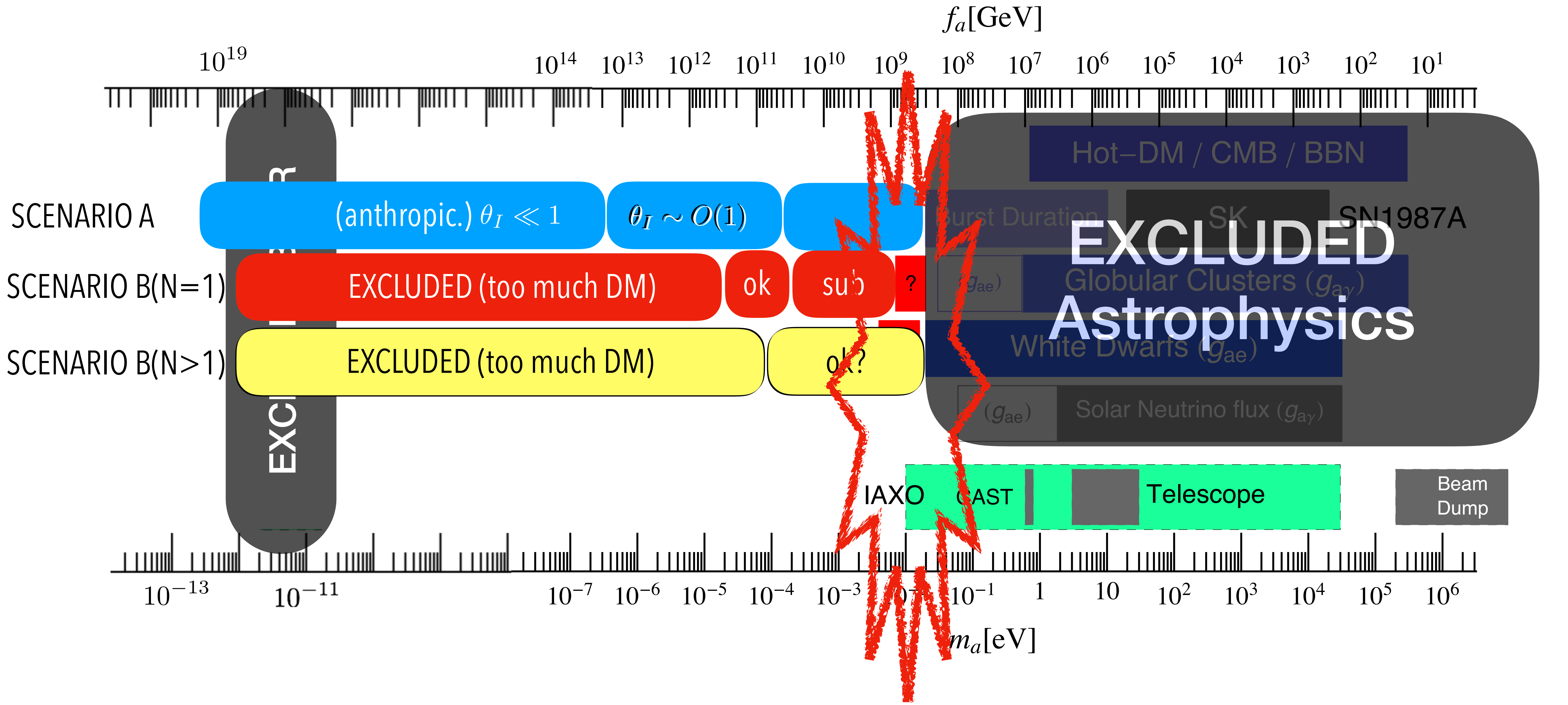


- 12 h tracking/day (sunset,dawn)
- 8 bores (60 cm diam)
- different Detectors
- dedicated X-ray optics
- Collaboration Formed 2017

In more detail

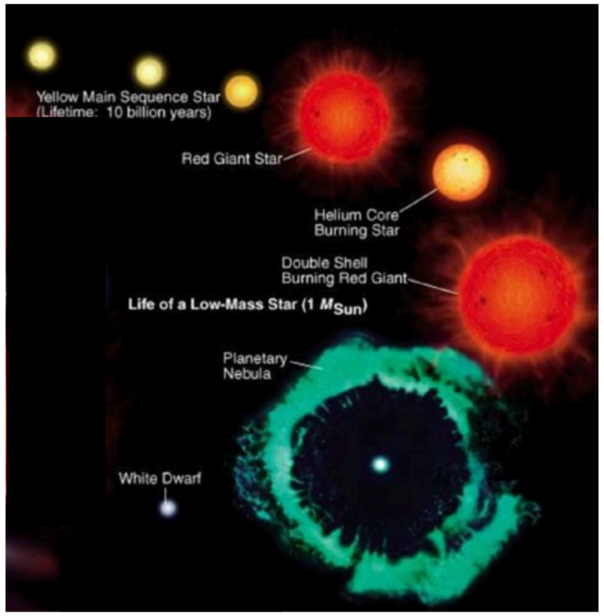


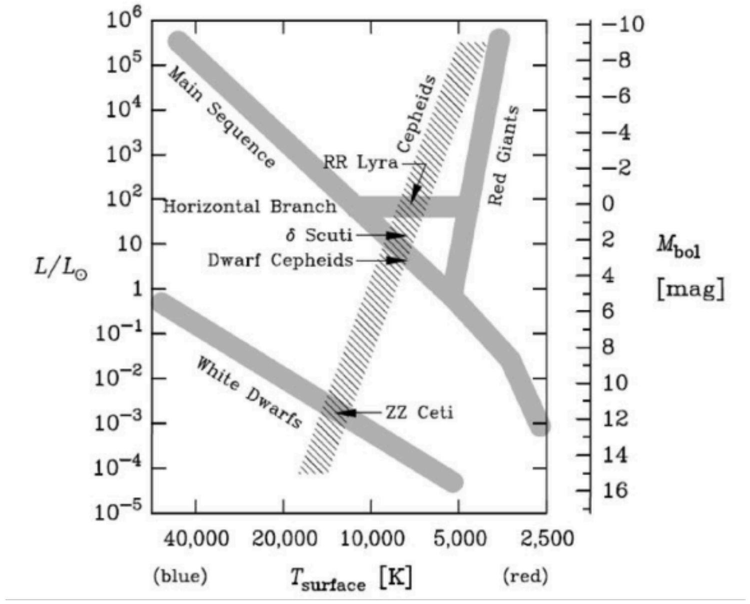
On the landscape



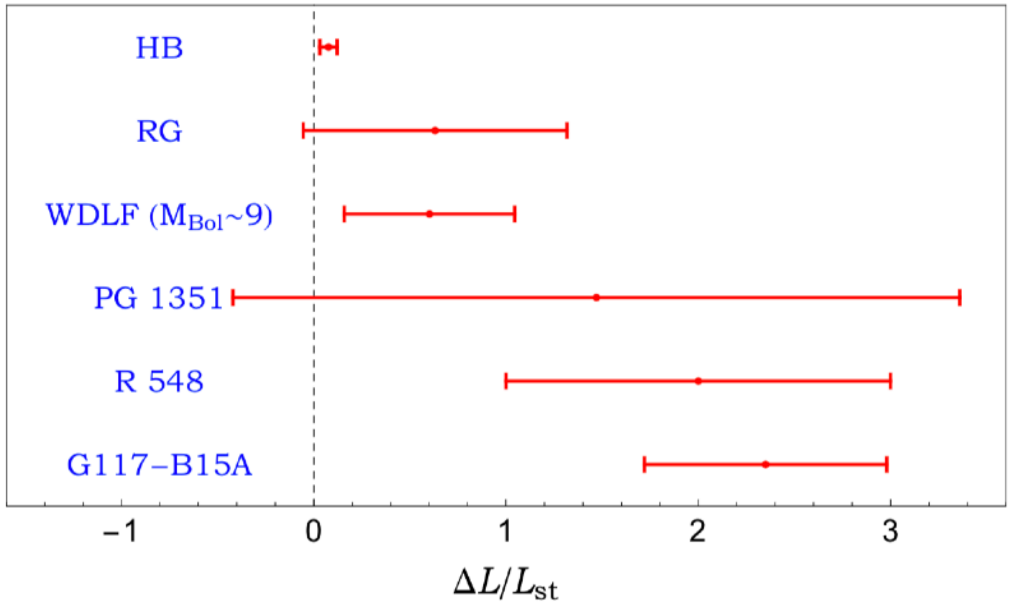
The meV frontier

- **Some stars seem to be cooling too fast; could be a sign for axion core emission!**



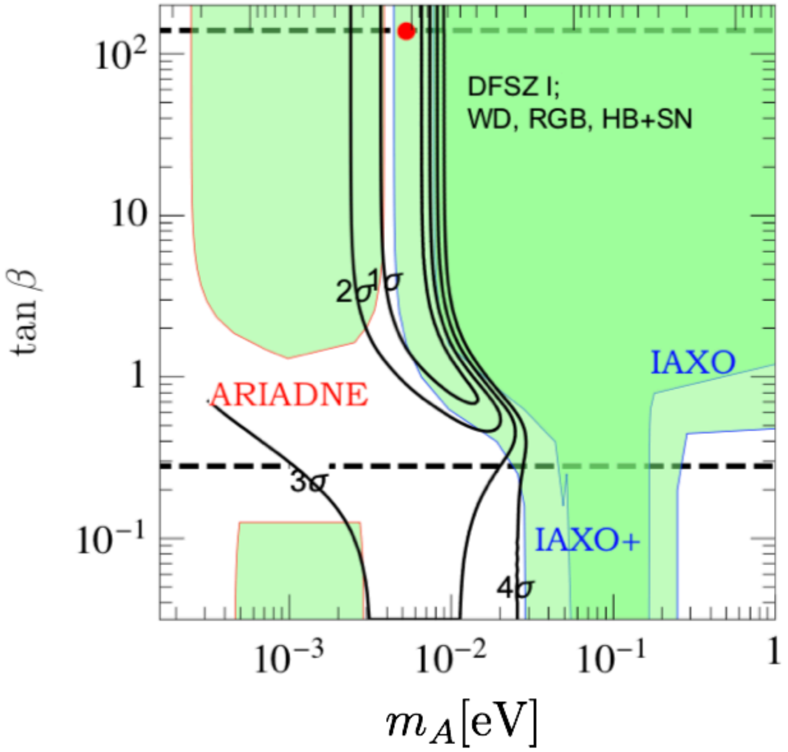


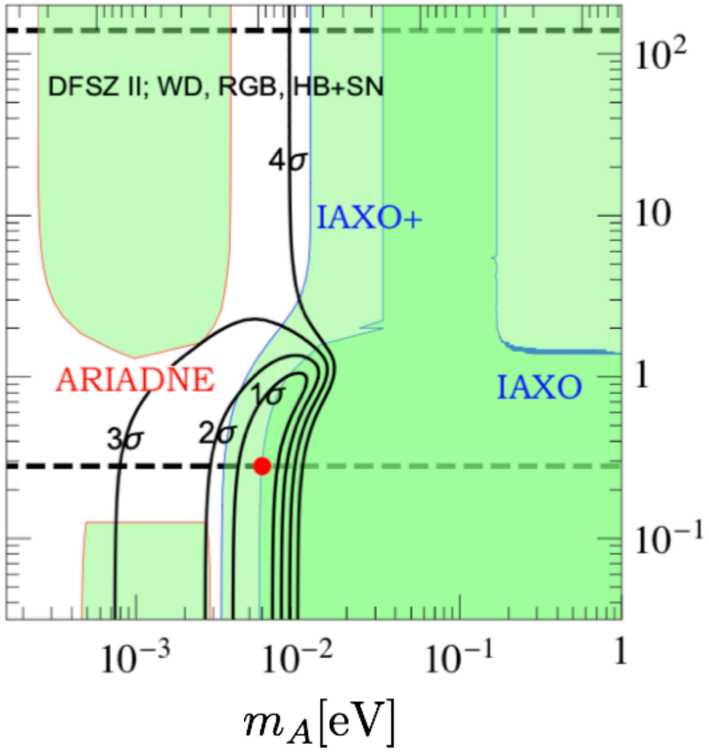
Giannotti 2015



- **Axion interpretation suggests axion-electron coupling (DFSZ?)**

Giannotti 2017





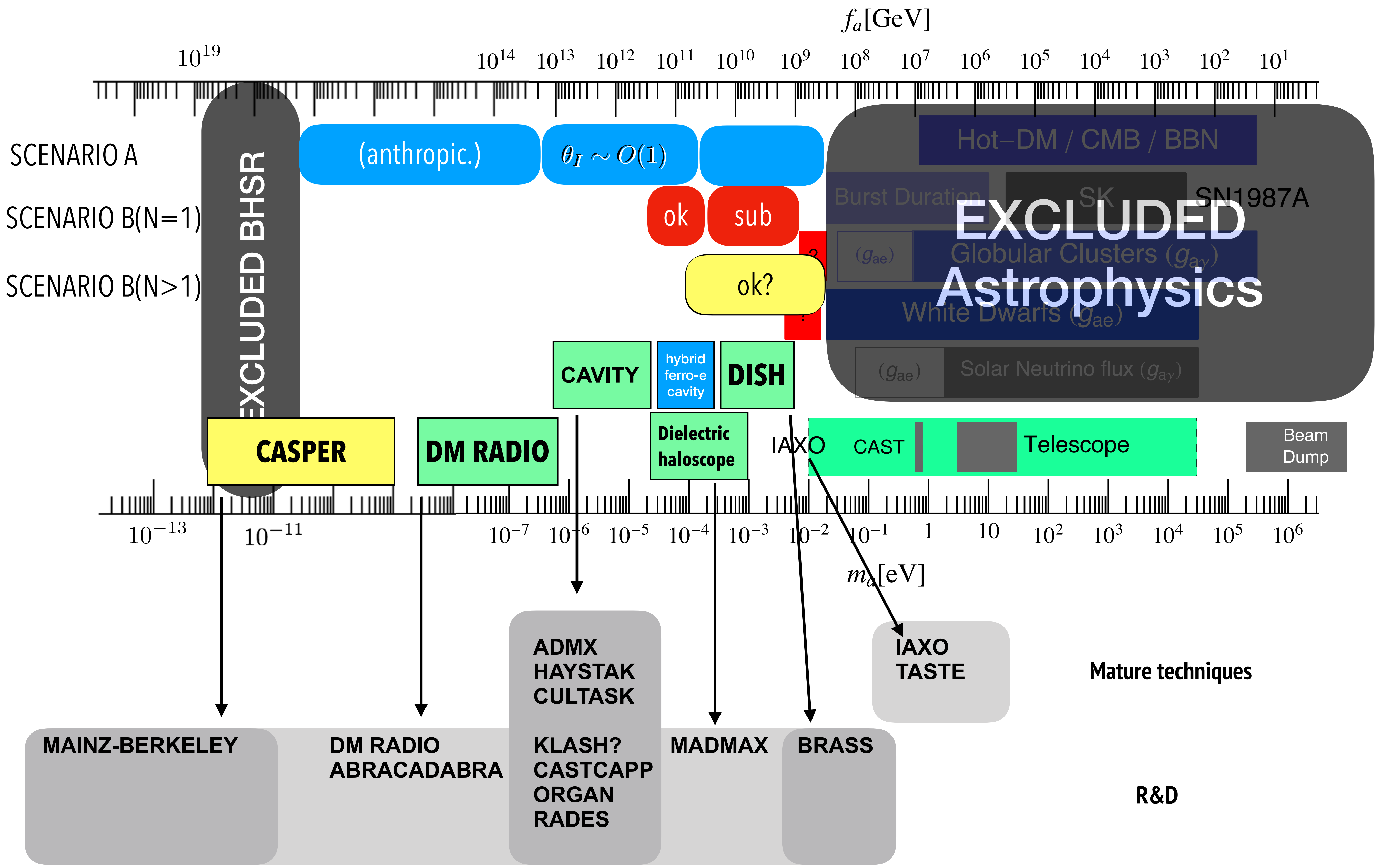
Isocontours show:

- *Best region fitting the anomaly (1,2,3,4s)*
- *Includes SN constraint !!!*
- *Dashed lines limit perturbativity of Yukawas*

Clear preference for m_A~6 meV

- *Detectable by IAXO!*

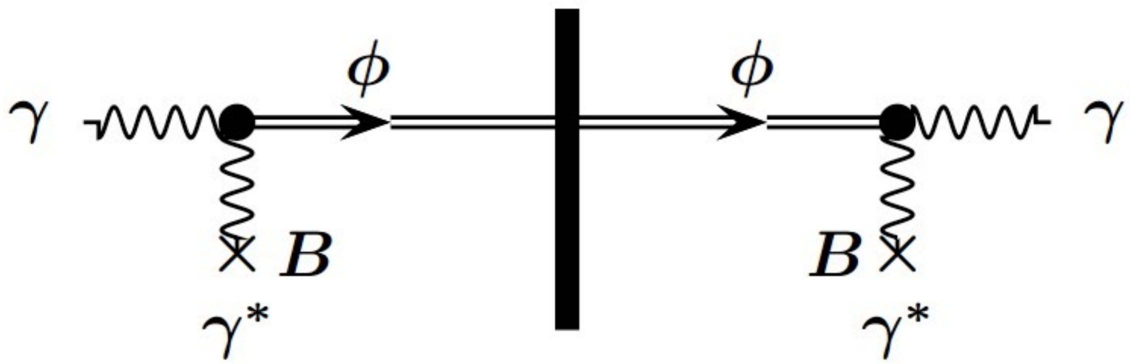
Experiments



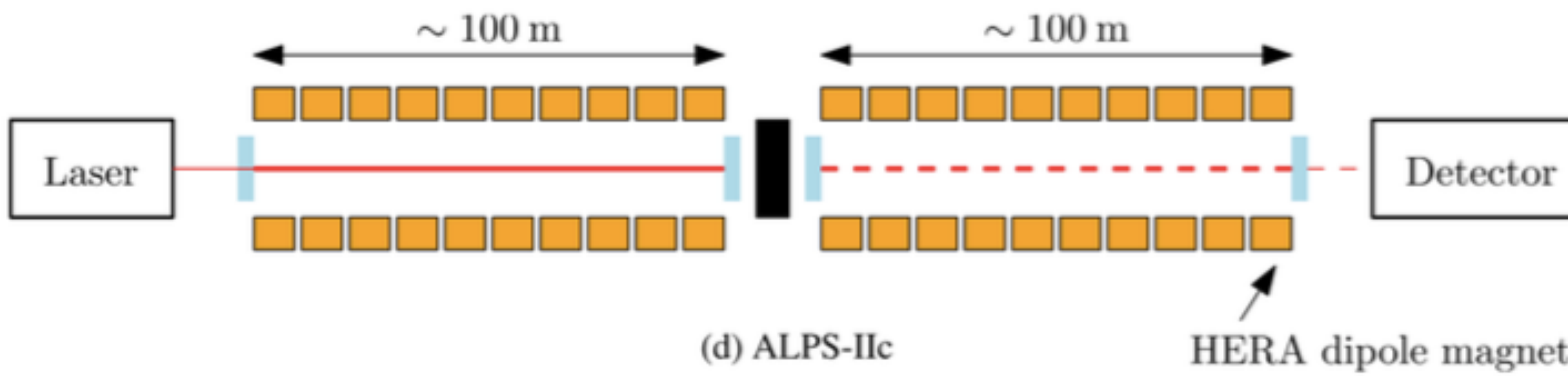
1.3 - Purely lab experiments

the ANY-Light-Particle-Search

Light shining through walls



Resonant regeneration in the receiving cavity



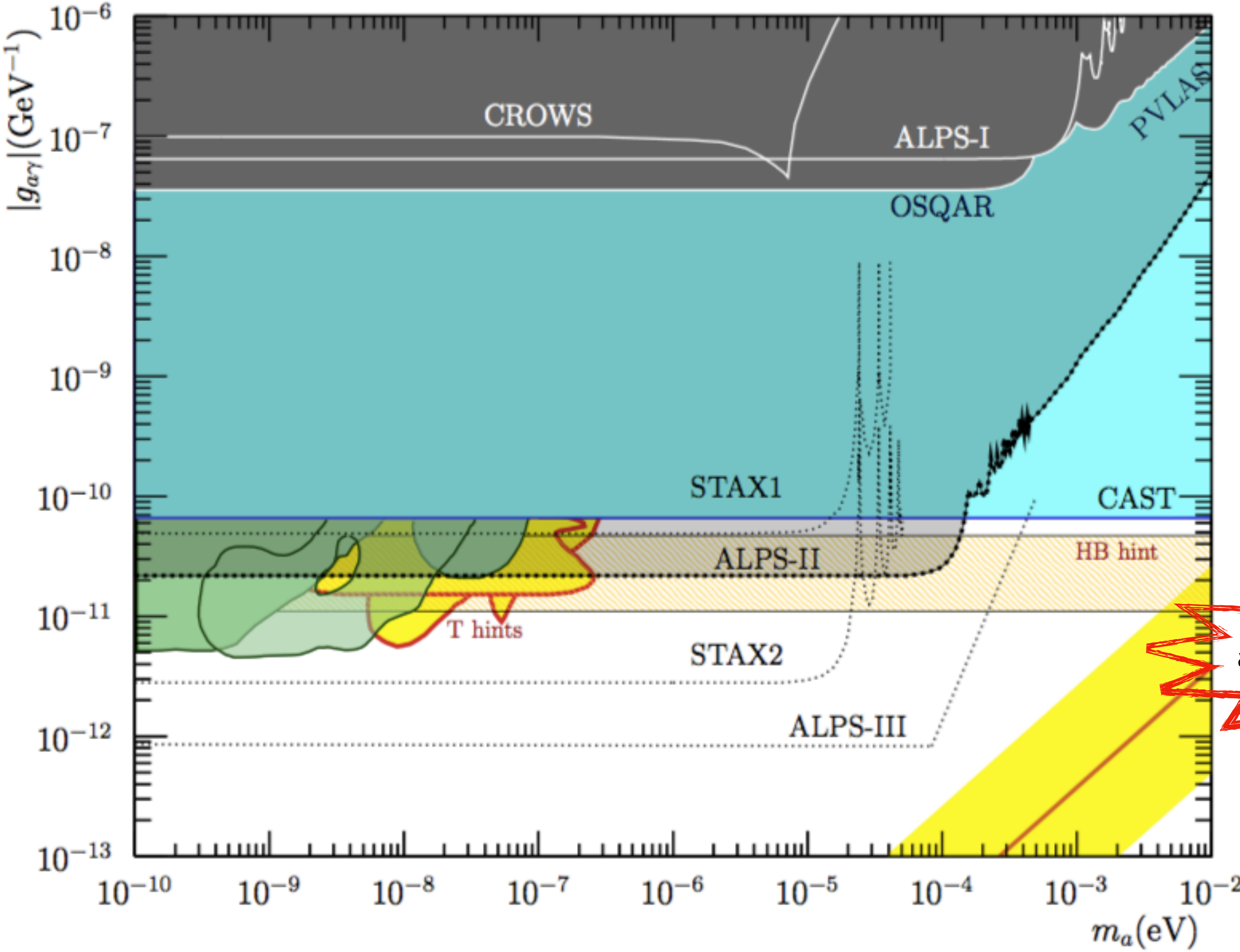
(d) ALPS-IIc

Exp.	Photon flux (1/s)	Photon E (eV)	B (T)	L (m)	B·L (Tm)	PB reg.cav.	Sens. (rel.)
ALPS I	$3.5 \cdot 10^{21}$	2.3	5.0	4.4	22	1	0.0003
ALPS II	$1 \cdot 10^{24}$	1.2	5.3	106	468	40,000	1
"ALPS III"	$3 \cdot 10^{25}$	1.2	13	400	5200	100,000	27

Experiment	status	B (T)	L (m)	Input power (W)	β_P	β_R	$g_{a\gamma}$[GeV^{-1}]
ALPS-I [427]	completed	5	4.3	4	300	1	5×10^{-8}
CROWS [429]	completed	3	0.15	50	10^4	10^4	9.9×10^{-8}(*)
OSQAR [428]	ongoing	9	14.3	18.5	-	-	3.5×10^{-8}
ALPS-II [430]	in preparation	5	100	30	5000	40000	2×10^{-11}
ALPS-III [431]	concept	13	426	200	12500	10^5	10^{-12}
STAX1 [432]	concept	15	0.5	10^5	10^4	-	5×10^{-11}
STAX2 [432]	concept	15	0.5	10^6	10^4	10^4	3×10^{-12}

STAX, ALPS III and beyond

ALPS with optical lasers, STAX with Microwave cavities … not so good for QCD…



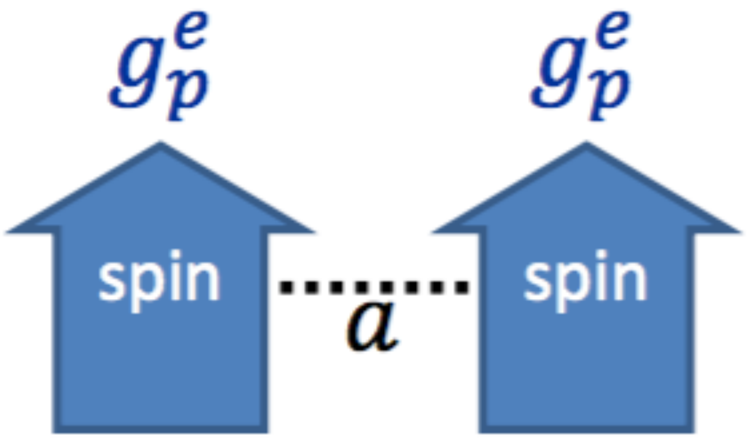
axion band $g_{A\gamma} \propto m_A$

Long-range forces

Wilzcek '84, Geraci 14

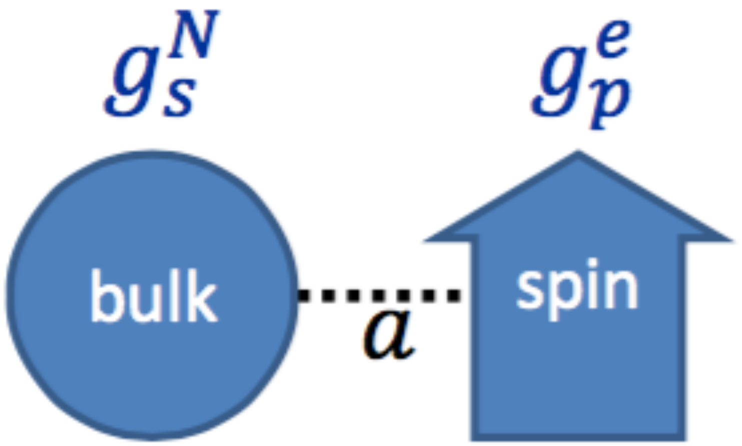
Long-range forces between macroscopic bodies

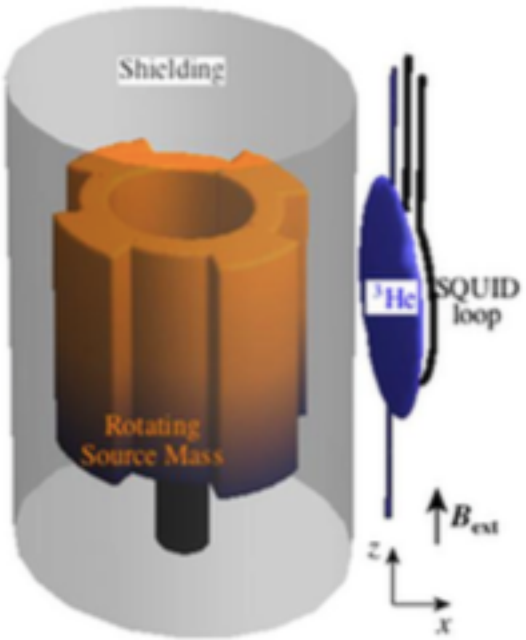
p-p forces are spin-spin ... very hard to measure!



In some case a tiny s-coupling can lead to a larger effect

s-p forces are number-spin ... much easier

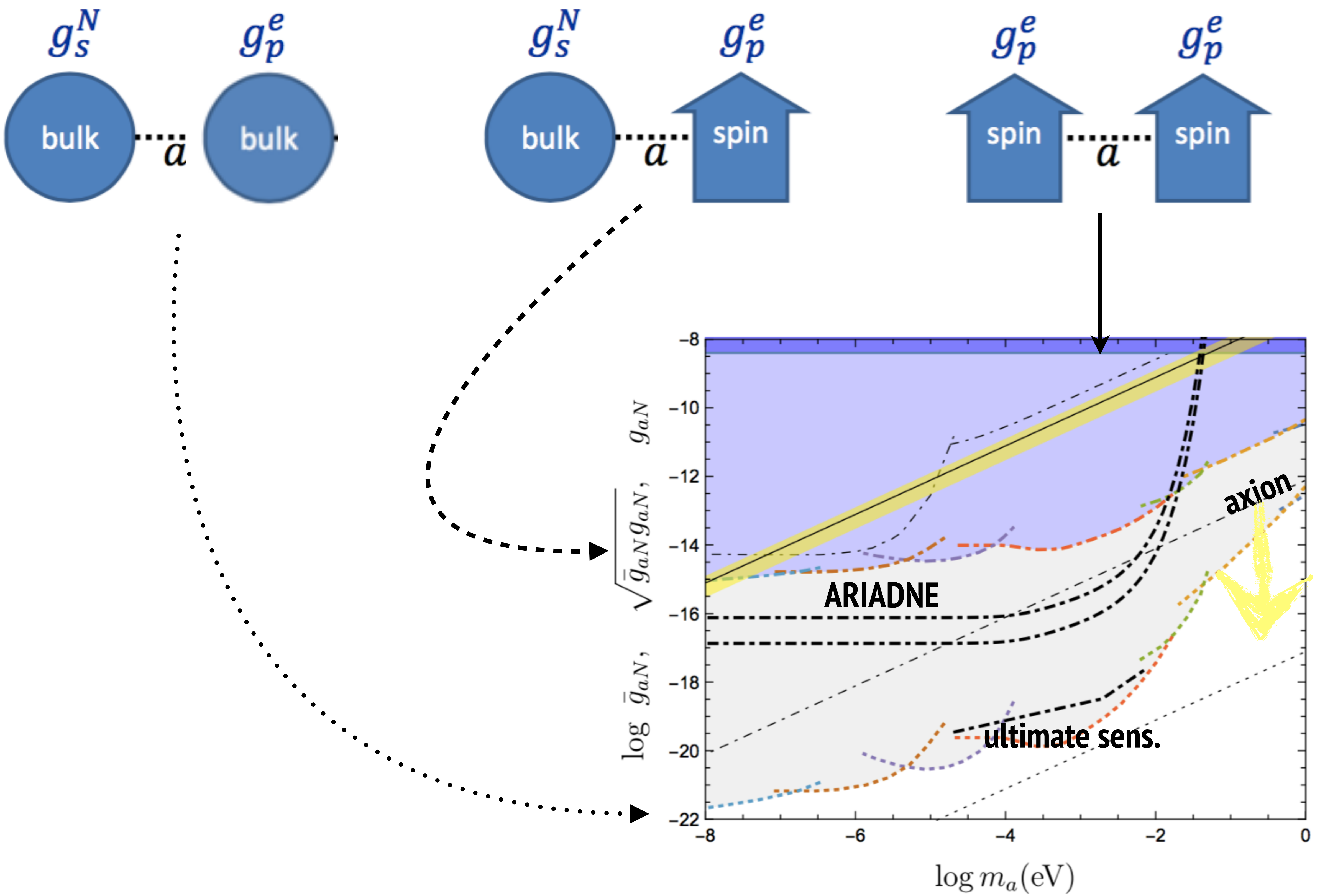




ARIADNE, University of Nevada in Reno

ARIADNE reach

Arvanitaki, Geraci 14



Flavoured axions

- Axions related to flavour/family symmetries induce Flavour violating decays

$$\Gamma(K^+ \to \pi^+ a) \simeq \frac{m_K}{64\pi} g^2_{aff'}$$

$$BR(\pi^+ a) < 7.3 \times 10^{-11} \quad (E787, E949)$$

(NA62, ORKA, KOTO improvement by ~ 70 on BR)

$$f_a \gtrsim \frac{\kappa_{sd}}{N} \times 7.5 \cdot 10^{10}\,\mathrm{GeV}\,,$$

model dependent coefficient

$$BR(B^+ \to K^+ a) < 10^{-8} \sim 10^{-6} \quad (Belle2?)$$

1 - Experiments